

Barium stars as tracers of s-process nucleosynthesis in AGB stars

I. 28 stars with independently derived AGB mass

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ABSTRACT

Context. Barium (Ba) stars are polluted by material enriched in the slow neutron capture (*s*-process) elements synthesised in the interior of their former asymptotic giant branch (AGB) companion star, which is now a white dwarf.

Aims. We aim to compare individual Ba star abundance patterns to AGB nucleosynthesis model predictions to verify if the AGB model mass is compatible with independently derived AGB mass, which was previously estimated using binary parameters and *Gaia* parallax data.

Methods. We selected a sample of 28 Ba stars for which both self-consistent spectroscopic observation and analysis were performed and, additionally, stellar mass determinations, via positioning the star on the Hertzsprung-Russell (HR) diagram and comparing with evolutionary tracks are available. For this sample of stars, we considered both previously (Y, Zr, Ce, and Nd) and recently derived (Rb, Sr, Nb, Mo, Ru, La, Sm, and Eu) elemental abundances. Then, we performed a detailed comparison of these *s*-process elemental abundances to different AGB nucleosynthesis models from the Monash and the FRUITY theoretical data sets. We simplified the binary mass transfer by calculating dilution factors to match the [Ce/Fe] value of each star when using different AGB nucleosynthesis models, and we then compared the diluted model abundances to the complete Ba-star abundance pattern.

Results. Our comparison confirms that low-mass (with initial masses roughly in the range 2–3 $M_{\odot}$), non-rotating AGB stellar models with ^{13}C as the main neutron source are the polluters of the vast majority of the considered Ba stars. Out of the 28 stars, in 21 cases the models are in good agreement with both the determined abundances and the independently derived AGB mass, although in 16 cases higher observed abundances of Nb, Ru, Mo, and/or Nd, Sm than predicted were present. For three stars, we obtain a match to the abundances only by considering models with masses lower than those independently determined. Finally, four stars show much higher first *s*-process peak abundance values than the model predictions, which may represent the signature of a physical (e.g. mixing) and/or nucleosynthetic process that is not represented in the set of models considered here.

Key words. stars: chemically peculiar – nuclear reactions, nucleosynthesis, abundances – stars: AGB and post-AGB

1. Introduction

Barium (Ba) stars are of undisputable significance for the understanding of nucleosynthetic processes in the interior of asymptotic giant branch (AGB) stars. A Ba star is a dwarf or a giant of F-K spectral class and typical metallicity from solar to roughly -0.6 dex (see, e.g. Fig. 3a in de Castro et al. 2016), whose spectrum show signatures of enhancement relative to solar of slow neutron-capture (*s* process) elements heavier than Fe, as first reported by Bidelman & Keenan (1951). Since the discovery of the binary nature of Ba stars (McClure et al. 1980; McClure 1983), it has become widely accepted that the overabundance of these elements originates from mass transfer from a former AGB (now white dwarf) companion.

Two different neutron source reactions can provide neutrons for the *s* process to produce the elements heavier than Fe in AGB

stars (Gallino et al. 1998; Lugaro et al. 2003; Cristallo et al. 2009; Käppeler et al. 2011). The bulk of the neutrons in low-mass AGB stars, in the range from $\approx 1.5 M_{\odot}$ (or maybe even lower; see discussion in Sect. 5 and Shetye et al. 2019) to $4 M_{\odot}$, are produced within the He-rich intershell region (i.e. the region located between the He- and the H-burning shells) in a so-called ^{13}C pocket through the $^{13}\text{C}(\alpha, n)^{16}\text{O}$ reaction. This reaction requires a lower temperature (≈ 100 MK) to be activated than the $^{22}\text{Ne}(\alpha, n)^{25}\text{Mg}$ reaction, which needs $T \gtrsim 300$ MK and is dominant instead in AGB stars of masses $\gtrsim 4 M_{\odot}$. There, temperatures can go above this limit during the recurrent convective episodes of He burning (the thermal pulses). The ^{22}Ne neutron source produces higher neutron densities than the ^{13}C neutron source, therefore activating branching points along the *s*-process path and leading, for example, to the overproduction of Rb with respect to neighbouring elements such as Sr and Zr.

As such, Rb has been used as an indicator of the AGB stellar mass (Abia et al. 2002; Karakas et al. 2012; van Raai et al. 2012; Roriz et al. 2021a).

To characterise s -process abundance patterns, ratios of the elements at the first (neutron magic number of 50, i.e. Sr, Y, Zr) and second s -process peaks (neutron magic number of 82; i.e. Ba, La, Ce) have been widely used. This is because the relative abundances of these elements can provide a direct constraint to the total number of free neutrons available at the s -process site (see, e.g. Cseh et al. 2018; Lugaro et al. 2020 for more details). Such ratios are particularly useful for Ba stars because the dilution due to the mass transfer affects all s -process elements in a similar way, and therefore their ratios closely reflect the intrinsically produced nucleosynthetic ratio.

In Cseh et al. (2018) we compared the large sample of 169 Ba stars observed by de Castro et al. (2016) to AGB models as a population and found that the distribution of the ratios of second-peak (Ce and Nd) to first-peak elements (Y and Zr) as a function of $[\text{Fe}/\text{H}]^1$ can be well described by non-rotating AGB models. We concluded that the metallicity is the best predictor of the s -process pattern, with lower metallicities favouring the production of the elements of the second peak relative to the first peak, as observed in the Ba stars. This results appear to be confirmed by recent observations of ten Ba stars (Shejeelammal et al. 2020). However, for each metallicity, a spread in the second to first-peak element ratios is present in the sample, which needs to be explained by variations in other stellar features, such as the initial mass, but also possibly rotation, and/or diffusive mixing. To disentangle these latter effects from the effect of the initial mass we need to compare the individual detailed stellar abundance pattern to single AGB models and verify if the mass of the stellar model corresponds to the mass of the AGB companion determined independently, when available.

To remove the effects of binary mass transfer and mixing on the secondary Ba star and select those effects directly related to nucleosynthesis, the observed $[\text{X}/\text{Fe}]$ abundance ratios should be compared to the models relative to each other, rather than in absolute terms. The mixing in our sample of giant Ba stars is mostly due to the fact that these stars undergo the first dredge up when ascending the first giant branch (Becker & Iben 1979; Boothroyd & Sackmann 1999), which mixes the material that was transferred onto the stellar surface with the whole envelope mass. To mimic the dilution due to the mass transfer and mixing, a free parameter is therefore introduced, usually referred to as the dilution factor (dil), representing the ratio between the total mass, that is the original envelope mass of the Ba star plus the mass transferred from the AGB star to its companion, and the transferred AGB material, and thus $\text{dil} = M_{\text{total}}/M_{\text{AGB,trans}} = (M_{\text{Ba,env}} + M_{\text{AGB,trans}})/M_{\text{AGB,trans}}$ (see, e.g. Husti et al. 2009; Bisterzo et al. 2012). For Ba dwarf stars without extended convective envelopes, the transferred s -process-rich material is deposited on top of the surface of the star and dilution is usually assumed to be inefficient², so dil is practically 1. In the case of giant Ba stars, such as those analysed by de Castro et al. (2016) and considered here, the effect of dilu-

tion cannot be neglected due to the large convective envelope in which the s -process enriched material is mixed after being transferred, and thus $\text{dil} > 1$.

Several studies have included comparison of individual AGB models and Ba star abundances including the estimate for calculations of dilution factors. For example, Husti et al. (2009) examined 34 Ba dwarfs and giants from two different studies and compared their abundance patterns to theoretical models calculated using evolutionary models from the FRANEC code (Chieffi & Straniero 1989) and the Torino s -process post-processing code. These authors found that AGB models in the mass range from $1.5 M_{\odot}$ to $3.0 M_{\odot}$ and dilution factors from 1 (for dwarf stars) to 30 were required to cover the observed abundances in the sample stars considered. A recent study of Shejeelammal et al. (2020) for Ba and CH stars provides abundances for many first and second s -process peak elements, and also included C, Rb, and Eu. Comparing the derived abundances to FRUITY nucleosynthesis models, these authors confirm that AGB stars with masses of $\lesssim 3 M_{\odot}$ are the typical former companions of Ba stars. However, for these previous studies only the Ba star metallicity ($[\text{Fe}/\text{H}]$) is available from the observations, while the initial AGB stellar mass is unknown and therefore can be treated as a free parameter to match the observations. This hampers the possibility of determining if variations or mismatches are due to a different stellar masses or to other stellar physics or nucleosynthetic features.

Thanks to the *Gaia* survey, it has recently become possible to provide better determinations for the masses of Ba stars (Jorissen et al. 2019) using the *Gaia* DR2 parallaxes (Gaia Collaboration 2018) and the spectral energy distribution of the stars. It is therefore possible to estimate the mass of the white dwarf companion using assumptions on the inclination of the binary system or the mass ratio of the members (or Q value, where $Q = M_{\text{WD}}^3/(M_{\text{Ba}} + M_{\text{WD}})^2$ – see Sect. 2 and Jorissen et al. (2019) – derived from the masses of the components, where M_{Ba} is the mass of the Ba star (Ba) and M_{WD} is that of the white dwarf (WD)). Then, from the WD mass one can obtain an estimate for the initial mass of the AGB star that polluted the observed Ba star (hereafter M_{AGB}^{J19}) using an initial-final mass relationship of El-Badry et al. (2018). By comparing the sample from Jorissen et al. (2019) to that of de Castro et al. (2016), we found 28 Ba stars in common, for which both an estimate of the initial mass of the AGB star and high-resolution spectroscopic abundances are available. Furthermore, the sample of de Castro et al. (2016) was analysed for more elements (Rb, Sr, Nb, Mo, Ru, La, Sm, and Eu, Roriz et al. 2021a,b) including a re-analysis of La, which was previously found to be affected by the saturation of strong lines (Cseh et al. 2018).

Armed with such a set of unprecedented constraints, we can now provide the best comparison for Ba star data to AGB models, to verify to which extent AGB model predictions match the data and to attempt to disentangle the potential effects of different stellar features on the s process. Furthermore, we provide estimates of the dilution factors, which, together with the known orbital parameters of these systems, will also pave the way to a better understanding of mass transfer in these systems, which is still uncertain (see, e.g. Sect. 5.4 of De Marco & Izzard 2017).

The paper is organised as follows. In Sect. 2, we present our sample data and summarise the new elemental abundances from Roriz et al. (2021a,b). In Sect. 3, we describe our methodology to find potential fits for each star using both the FRUITY and the Monash set of yields and to calculate the dilution factor for each potential solution. In Sect. 4, we present our comparison between models and determined abundances for each individual

¹ Throughout the paper we will use the usual spectroscopic notation $[\text{Y}/\text{X}] = \log_{10}(\text{Y}/\text{X})_{\text{star}} - \log_{10}(\text{Y}/\text{X})_{\odot}$.

² We note that mixing mechanism as thermohaline mixing could still play a role. See, for example, Stancliffe et al. (2007) and Stancliffe & Glebbeek (2008) for the case of carbon-enhanced metal-poor (CEMP) stars (the low-metallicity counterparts of Ba stars) and Husti et al. (2009) for Ba stars. Although, Denissenkov & Pinsonneault (2008) and Aoki et al. (2008) suggest that thermohaline mixing is inefficient in dwarf CEMP stars.

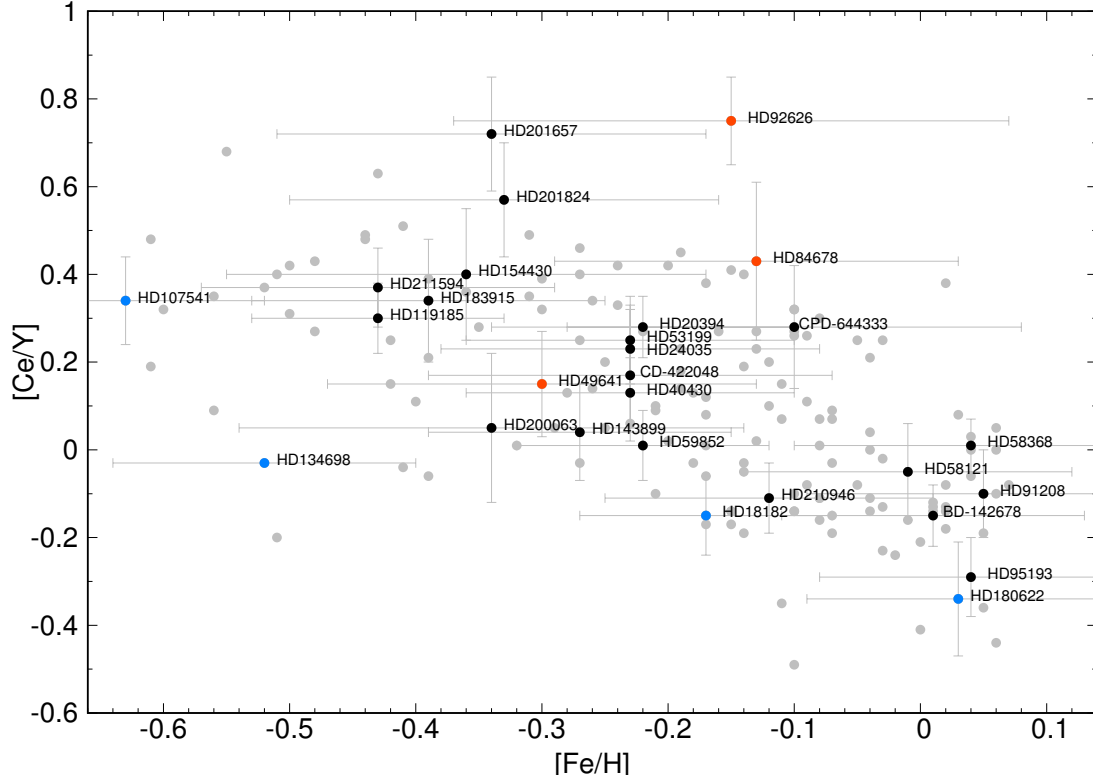


Fig. 1. [Ce/Y] values as function of metallicity for the stars analysed in de Castro et al. (2016) and Cseh et al. (2018) (grey dots), with the 28 stars analysed in this work indicated by the names and highlighted as black, blue, and red dots, with colour depending on Groups 1, 2, and 3 presented in Sect. 4, respectively. For sake of clarity, the error bars are marked only for the stars analysed in this paper, and the three Ba stars at $[\text{Fe}/\text{H}] \approx -1$ are not shown.

star. Finally, we discuss our results in Sect. 5 and summarise them in Sect. 6.

2. Sample stars and abundances

2.1. Sample stars

Out of the 169 Ba giants analysed in de Castro et al. (2016) and Cseh et al. (2018), 28 stars were also included in the sample of Jorissen et al. (2019), which was aimed at determining the orbital parameters of Ba and related stars. The locations in the [Ce/Y] versus [Fe/H] plane of the stars analysed in this paper relative to the whole sample of Cseh et al. (2018) are shown in Fig. 1, which shows that the [Ce/Y] ratio of the 28 selected stars is typical of the entire sample, and it covers the whole range of the sample stars. This means that our sample allows us to investigate the origin of the spread present at any given [Fe/H]. The masses of the Ba stars (M_{Ba}) were determined by Jorissen et al. (2019) from the location of these stars on the HR diagram and using the STAREVOL evolutionary models (Siess et al. 2000). Jorissen et al. (2019) also calculated the companion white dwarf (WD) masses (M_{WD}) from the mass function of the system assuming a random distribution of the orbital inclination and assuming that $Q = M_{\text{WD}}^3 / (M_{\text{Ba}} + M_{\text{WD}})^2$ is a constant (the ‘constant Q assumption’). The initial AGB masses ($M_{\text{AGB}}^{\text{J19}}$) were then determined from the estimated WD masses using the initial-final mass relation derived by El-Badry et al. (2018), as described in Jorissen et al. (2019); see their Fig. 14, and assuming that the binary evolution did not affect this relation (for more details on this method, see also Escorza et al. 2017; Van der Swaelmen et al. 2017).

Many uncertainties affect the derivation of this initial AGB mass. First, in relation to the mass of the Ba star, the *Gaia* DR2 parallaxes (Gaia Collaboration 2018) were derived using single-star solutions. Since our sample stars are in a binary system, this might already introduce errors in the derivation of M_{Ba} . As for the re-normalised unit-weight error (RUWE) – a value when above roughly 1–1.4 indicates potential problems with the single-star astrometric solution (see, e.g. El-Badry et al. 2021; Stassun & Torres 2021) – only one star in our sample (BD-14°2678) has $\text{RUWE} < 1$ (0.919), while 18 stars have $\text{RUWE} > 1.4$ (Gaia Collaboration 2020). Second, there are several problems in deriving the AGB initial mass from the Ba stars mass. On top of the strong constant Q assumption and the assumption that the orbital inclination is randomly distributed, neglecting the binary evolution in the initial-final mass relationship can introduce further uncertainties. In spite of these difficulties, baseline evidence of the validity of the method is provided by the fact that, as shown by Jorissen et al. (2019), most of the systems analysed have $M_{\text{AGB}}^{\text{J19}} / M_{\text{Ba}} > 1$, as expected. Overall, since the initial-final mass relationship of El-Badry et al. (2018) gives an initial AGB mass of $2.75 \pm_{0.31}^{0.36}$ for a $0.67 M_{\odot}$ WD, and higher mass WDs have larger uncertainty, we assume that the uncertainty on the derived initial AGB mass is of at least $\pm 0.5 M_{\odot}$.

Finally, we note that the classification criteria for Ba stars is still not homogeneous in the literature. In de Castro et al. (2016), a star is considered a Ba star if the condition $[s/\text{Fe}] \geq 0.25$ dex is fulfilled, where s is the average of Y, Zr, La, Ce, and Nd. In Jorissen et al. (2019), the [Ce/Fe] ratio was used to classify and distinguish between strong and mild Ba stars: if the value is over ≈ 1 dex, a Ba star is classified as strong, while mild Ba stars have

[La/Fe] and [Ce/Fe] values between 0.2 and 1 dex. In the next sub-section, we categorise the sample stars into three groups, depending on the match between the s -process elemental abundances and the AGB models. Stars with good agreement between the models and the s -process peak elements belong to Group 1, Group 2 includes four stars with low estimated Ba and AGB mass ($< 2 M_{\odot}$) and with higher first s -process peak than typically predicted, while Group 3 includes the three stars in our sample that have the highest initial AGB mass ($> 3.8 M_{\odot}$), as derived under the constant Q assumption (Jorissen et al. 2019).

2.2. Elemental abundances

On top of Y, Zr, Ce, and Nd from de Castro et al. (2016), our present analysis has the advantage of determining new elemental abundances for Rb, Sr, Nb, Mo, Ru, La, Sm, and Eu, which allow us a more detailed comparison between models and observational data. Rb abundances were published by Roriz et al. (2021a) for the whole sample of de Castro et al. (2016), while abundances for the other species are presented in Roriz et al. (2021b). As mentioned in Sect. 1, Rb is a key element to distinguish between reactions responsible for the neutron source in the polluting AGB star and also for the determination of the AGB mass, since its abundance depends on the activation of branching points on the s -process path.

The final abundances were computed using the current version of the spectral analysis code MOOG (Snedden 1973), which assumes a plane-parallel atmosphere model, under the conditions of local thermodynamic equilibrium (LTE), from either equivalent width measurements or spectral synthesis of the atomic lines considered. For the [X/Fe] ratios considered in this study, uncertainties were estimated from the variation in the abundances due to changes in the parameters of the atmospheric models, as well as from dispersion of the abundance, when three or more lines were available for a given element. We invite the reader to consult de Castro et al. (2016) and Roriz et al. (2021a,b) for details of the analysis and the list of lines used to obtain the abundances considered in this work.

Although [La/Fe] ratios were published in de Castro et al. (2016), some of the stars showed unexpectedly high values when compared to the ratios [Ce/Fe] and [Nd/Fe], as pointed out in Cseh et al. (2018). Since La, Ce, and Nd belong to the same s -process peak, these elements are expected to exhibit similar overabundances relative to solar. This is, for example, the case of the star CPD −64°4333, for which de Castro et al. (2016) reported [La/Fe] = 2.52 dex, approximately 1 dex higher than the values of [Ce/Fe] and [Nd/Fe] for the same object. Therefore, La abundances were re-determined (Roriz et al. 2021b). We compare the new [La/Fe] ratios to the [Ce/Fe] ratios in Fig. 2 for the sample considered in the present study. The new [La/Fe] ratios are consistent with the [Ce/Fe] ratios. In Fig. 3, we show the [X/Fe] ratios versus metallicity for elements between Rb and Eu, highlighting the stars belonging to our Group 1, 2, and 3 classifications (see Sect. 4).

3. Methods

3.1. Stellar models

We compared the abundances of each individual Ba star with the two most extensive currently available set of AGB nucleosynthesis models: Monash Lugaro et al. (2012), Fishlock et al. (2014), Karakas & Lugaro (2016), Karakas et al. (2018) and FRUITY Cristallo et al. (2009, 2011, 2015a, 2016), Piersanti et al. (2013).

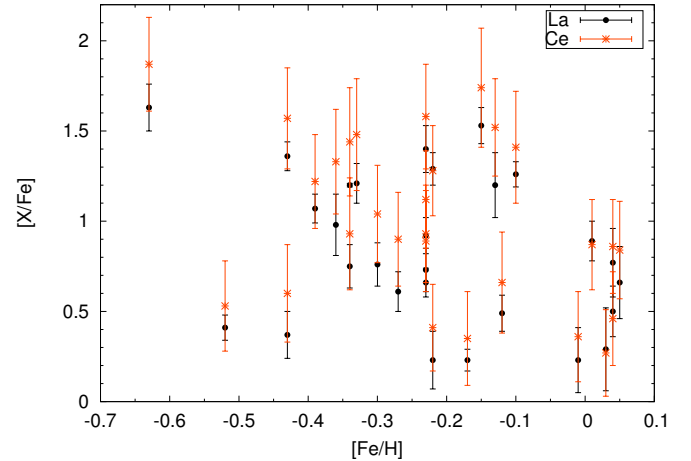


Fig. 2. [Ce/Fe] (orange asterisks) and the new [La/Fe] (black dots) abundances for our sample stars.

Other sets of models were not taken into account because of the limited number of models in mass and metallicity. Most of the Monash models with a given mass and metallicity were also calculated when varying the masses of the partial mixing zone (M_{PMZ}), leading to the formation of the ^{13}C pocket different from the standard value reported in Karakas & Lugaro (2016). This parameter mostly affects the absolute amount of the produced s -process material (see also Buntain et al. 2017). If more models with different pocket sizes are applicable to a star, we show all of them in our comparison. New unpublished Monash models with masses of 2 and $3 M_{\odot}$ were calculated for $Z = 0.01$ ([Fe/H] = −0.15) and with standard $M_{\text{PMZ}} = 2 \times 10^{-3} M_{\odot}$. This was done with the aim of having more Monash models closer in metallicity to several stars in our sample.

Among the FRUITY models, the largest s -process production occurs in the models with an extended ^{13}C pocket (labelled as ‘TAIL’). Furthermore, FRUITY models that include rotation with different initial rotation velocities were also considered here. As was shown by Cseh et al. (2018), rotating models have significantly lower values of the [Ce/Y] ratio than the large Ba star sample of de Castro et al. (2016). Nevertheless, for the sake of completeness we also included the FRUITY rotating models when applicable as a match for a given Ba star. We note that we only selected models with $[s/\text{Fe}] \geq 0.25$ dex (where s is the average of Y, Zr, La, Ce, and Nd), since this criteria was applied as a limit for our sample of Ba stars.

The models we used to compare to each individual star have metallicity [Fe/H] in the range of the uncertainties of the given star, and initial mass close to that estimated by Jorissen et al. (2019) for the same star. In some cases, we extended the range of mass of the considered models if the initial search for a good match with the derived [Rb/Fe] that, as described above, is the strongest diagnostic of stellar mass was not successful. For some stars (the 3 stars in Group 3 described in Sect. 4.3, specifically), there was no matching model found nearby the estimated values. In those cases, we considered all the models in order to find the most appropriate match, even if they were far from the mass estimates by Jorissen et al. (2019).

In Fig. 4, we show an example of solar metallicity ([Fe/H] = 0) Monash and FRUITY models between 1.5 and $4 M_{\odot}$ with their standard ^{13}C pockets. The top panel of the figure shows that the FRUITY models have typically lower abundances, which is likely due to smaller third dredge-up (TDU)

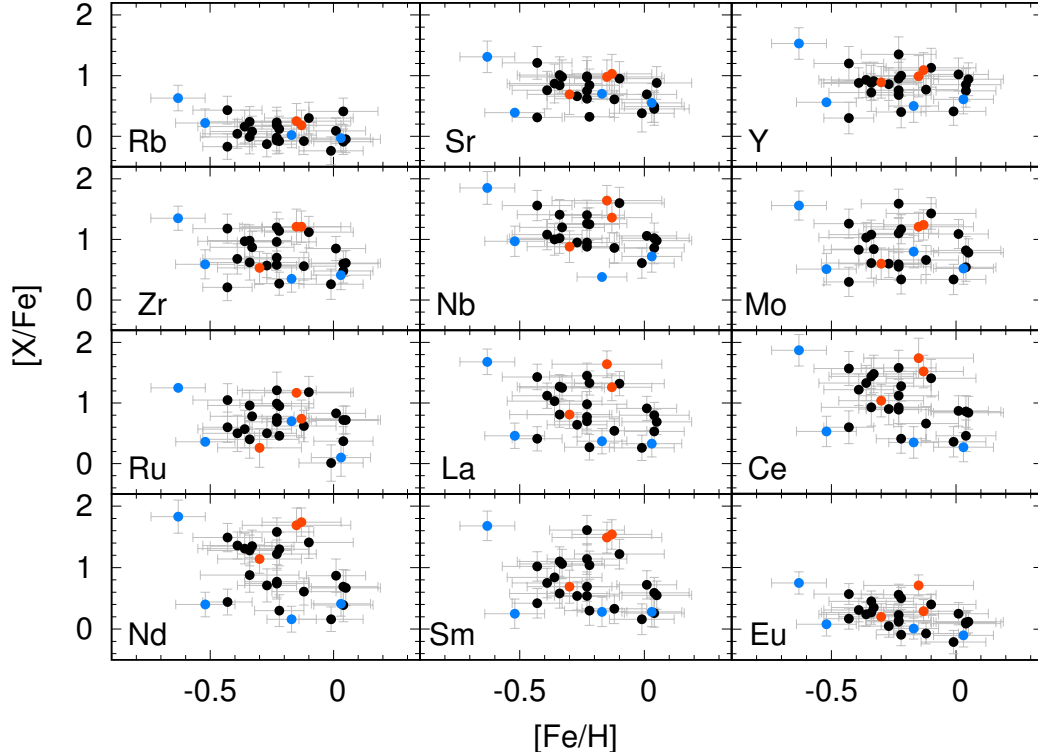


Fig. 3. Heavy element abundances for all the sample stars. The colours are the same as in Fig. 1.

efficiencies and smaller ^{13}C -pocket masses, on average. In fact, the Monash models typically experience more thermal pulses (TP) and show a deeper TDU, along with hotter burning at the base of the TP and of the envelope. The largest difference is between the $4 M_{\odot}$ models. This can be attributed to the fact that the ^{22}Ne neutron source is more active in the Monash models than in the FRUITY from mass $\approx 4 M_{\odot}$. Also, the decrease in the TDU efficiency has a large impact on the FRUITY $4 M_{\odot}$ and higher mass models, strongly reducing the final surface abundances of the s -process elements compared to lower mass models (Cristallo et al. 2015a). The $4 M_{\odot}$ Monash model still has an active ^{13}C neutron source, thus producing more s -process elements compared to the FRUITY models. For a detailed comparison of the two model sets in relation to the production of s -process elements, we refer the reader to Sect. 5 of Karakas & Lugaro (2016). Nevertheless, the relative abundance distributions are close to each other with similar height of the first two s -process peaks, and a lower third process peak. Rb is the element most sensitive to the stellar mass, with an abundance that increases with stellar mass, relatively to the first peak elements.

Finally, we note that other stellar processes not explicitly considered in the models used here could affect our analysis. For example, the models used in this study do not include any core overshoot (Cristallo et al. 2011, 2015b; Karakas & Lattanzio 2014). Core overshoot would increase the mass of the H-exhausted core at the end of the main sequence, and thus the star would behave as if it started with a slightly higher mass. The TDU may therefore occur at lower initial stellar masses than in the models used here leading to degeneracy between different masses. From studies of the core He-flash (e.g. Bertelli et al. 1986), the maximum mass shift is from $\approx 2 M_{\odot}$ to $\approx 1.8 M_{\odot}$ (roughly $0.2 M_{\odot}$ lower). This means that our models at around $1.5 M_{\odot}$ would be lowered to $\approx 1.3 M_{\odot}$ when including the core overshoot.

3.2. The effect of dilution

The effect of dilution after the mass transfer from the former AGB companion cannot be neglected, since all of the sample stars analysed in this paper are giant stars with an extended convective envelope. This means that the transferred material was mixed with the envelope of the current Ba star because of the first dredge up, lowering the abundances relative to Fe as compared to their value in the envelope of the polluting AGB star. The diluted model abundances ($[\text{X}/\text{Fe}]_{\text{dil}}$) were calculated as follows:

$$[\text{X}/\text{Fe}]_{\text{dil}} = \log \left(10^{[\text{X}/\text{Fe}]_{\text{ini}}} \times (1 - \delta) + 10^{[\text{X}/\text{Fe}]_{\text{AGB}}} \times \delta \right), \quad (1)$$

where $[\text{X}/\text{Fe}]_{\text{ini}}$ is the initial (solar) abundance ratio assumed in all the models and corresponds to the initial composition of the Ba star envelope, $[\text{X}/\text{Fe}]_{\text{AGB}}$ is the final surface abundance ratio predicted by the AGB models, and $\delta = 1/\text{dil}$. We derived the value of δ for each star and each model by equating $[\text{Ce}/\text{Fe}]_{\text{dil}}$ to the $[\text{Ce}/\text{Fe}]$ ratio observed in the Ba star. We subsequently applied this same value of δ to all the other elemental abundances calculated with Eq. (1) above.

In the bottom panel of Fig. 4, we apply a dilution to normalise all the models to the same $[\text{Ce}/\text{Fe}]$ value. This representation even more clearly shows the changes in $[\text{Rb}/\text{Fe}]$, along with variations of the first peak elements with varying the mass. As we move from lower to higher mass models, $[\text{Rb}/\text{Fe}]$ increases with increasing mass, for similar δ values. Generally, FRUITY models have higher δ values than Monash models, and this is due to the lower predicted abundances at the AGB surface.

Since $\delta = 1$ means that the AGB material is deposited on top of the Ba star without any mixing to allow some mixing of the transferred material with the envelope of the giant, we only selected models with $\delta \leq 0.9$ as possible solutions. This value would imply that 90% of the whole Ba star envelope is made of AGB material, which is still unrealistic but allows us to

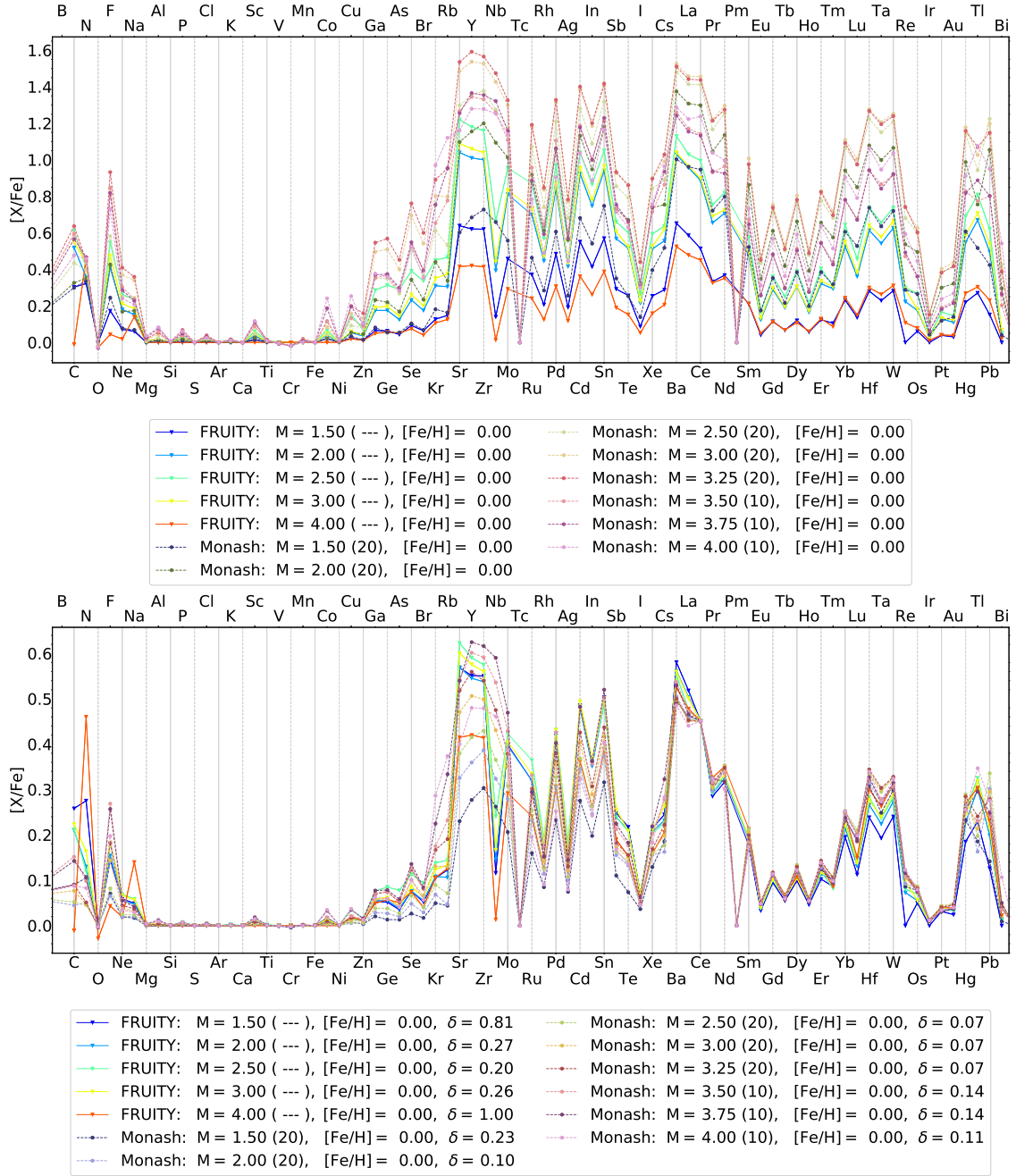


Fig. 4. Example of selected models with standard ^{13}C pocket showing the final surface abundances without applying any dilution ($\delta = 1$ for all models, *top panel*) and by applying a dilution (*bottom panel*) such that $[\text{Ce}/\text{Fe}] = 0.451$ dex, which is the value produced by the FRUITY $4 M_{\odot}$ model. Thus, this model has $\delta = 1$. Specific properties of the FRUITY models are given in parentheses after the mass of each model: ‘---’ indicates the standard models, while ‘TAIL’, ‘rotating’ and ‘rotating TAIL’ models with 60 km s^{-1} initial rotational velocity are indicated by (ext), (rNN), and (T60), respectively, where NN is the initial rotational velocity in km s^{-1} . The numbers given in parentheses after the mass of each Monash model represent the M_{PMZ} in units of $10^{-4} M_{\odot}$. We note that the high N and Na values shown by the undiluted $4 M_{\odot}$ FRUITY model are due to production by the first dredge up, the other models do not show these high values because they are diluted. We also note that the FRUITY Nb and Sm abundances are lower than in the Monash models because ^{93}Zr and ^{147}Pm are not decayed.

investigate the behaviour of the δ values. We discuss this point in more detail in Sect. 5.3.

3.3. Final output

Following the selection of the model and the dilution, our algorithm automatically generates a figure for each star. This shows, as function of the atomic number, $[X/Fe]$ for all the available observational data (from the light to the heavy elements) and

$[X/Fe]_{\text{dil}}$ from all the models of the selected mass that match the first peak elements, especially Rb. The reason behind the choice of matching these elements is that we want to verify if the independent estimate of M_{AGB}^{J19} is consistent with our mass estimate based on the chemical pattern. Since we are guided by the observation of the first peak elements only, we decided to not use the χ^2 method usually employed to find the best fitting model to the derived abundance pattern. Furthermore, in our case because of the fine grid of models that we used for the comparison, in

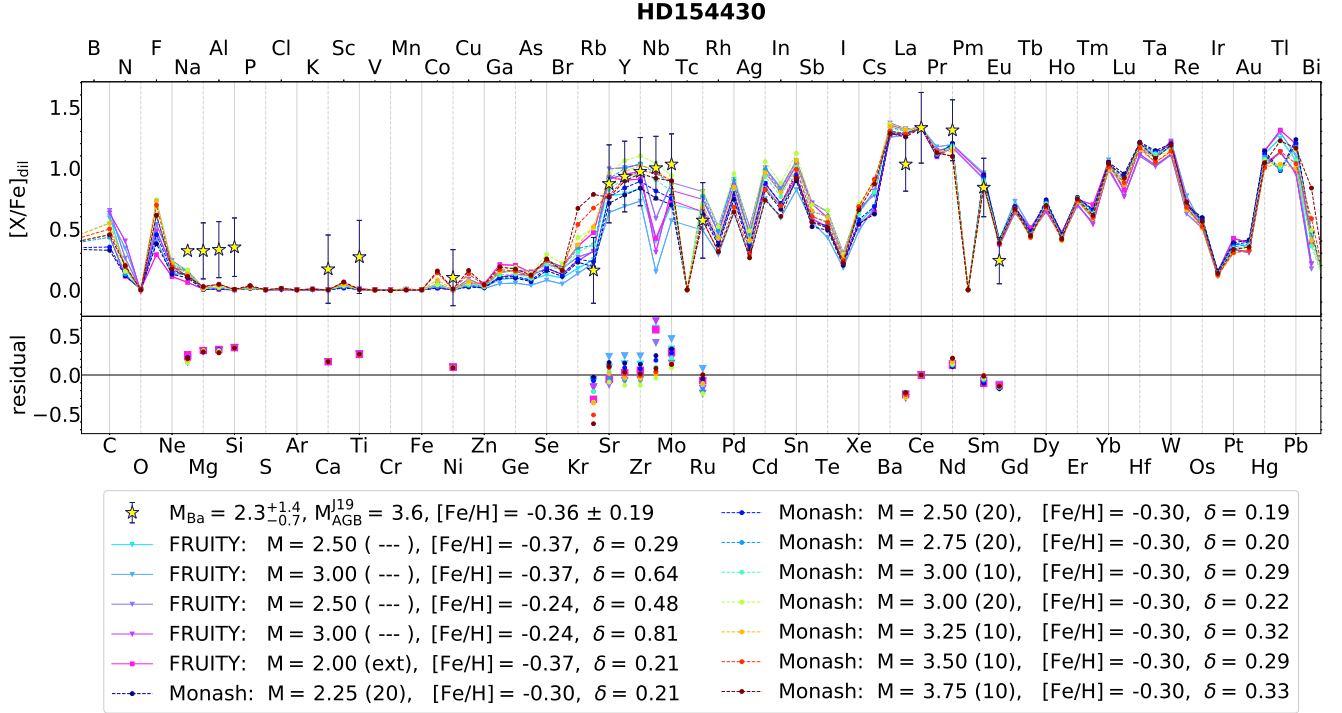


Fig. 5. Available elemental abundances from [de Castro et al. \(2016\)](#) and [Roriz et al. \(2021a,b\)](#) and a selection of AGB models. *Top panel:* comparison of the abundance pattern of HD 154430 and the final surface abundances of different AGB models after applying dilution. The residuals for each model are shown in the *bottom panel*. The label box lists the mass of the current Ba star (M_{Ba}), the estimated initial mass of the AGB component (M_{AGB}^{J19}) from [Jorissen et al. \(2019\)](#), and the parameters of the plotted models. All the masses are given in $M_{\odot}$ units. Specific properties of the FRUITY models are given in parentheses after the mass of each model: ‘---’ indicates the standard models, while ‘TAIL’, ‘rotating’, and ‘rotating TAIL’ models with 60 km s^{-1} initial rotational velocity are indicated by (ext), (rNN), and (T60), respectively, where NN is the initial rotational velocity in km s^{-1} . The numbers given in parentheses after the mass of each Monash model represent the M_{PMZ} given in units of $10^{-4} M_{\odot}$.

terms of mass, metallicity, and ^{13}C pocket size (see Tables B.1 and B.2 for FRUITY and Monash models, respectively), and the similarities between many of them (see Fig. 4), there would be many degenerate solutions with similar χ^2 values, and we would not obtain any strong indication. We can see this, for example, in relation to the case of HD 95193 (Fig. A.9), where 12 out of the 14 models plotted, of masses between 2.0 and $3.5 M_{\odot}$, give practically the same results, especially for the Monash models. For these reasons, we did not use the χ^2 method in this study and preferred to first rely on the Rb abundance to independently verify the initial AGB mass of the sample stars.

The residual difference $[X/Fe] - [X/Fe]_{\text{dil}}$ is plotted in the bottom panel of each comparison figure. Below the figure, the legend lists the mass of the Ba star (M_{Ba}) and the initial mass of the AGB star (M_{AGB}^{J19}) in $M_{\odot}$ units from [Jorissen et al. \(2019\)](#) and the $[\text{Fe}/\text{H}]$ from [de Castro et al. \(2016\)](#). All the plotted models with their corresponding line type in the figure and δ value used are also indicated by their set (Monash or FRUITY), mass M in $M_{\odot}$, M_{PMZ} in case of the Monash models (given in parentheses after the mass of each model in units of $10^{-4} M_{\odot}$) and metallicity $[\text{Fe}/\text{H}]$. For the FRUITY models the ‘TAIL’, ‘rotating’ and ‘rotating TAIL’ models with 60 km s^{-1} initial rotational velocity are indicated by (ext), (rNN), and (T60), respectively, where NN is the initial rotational velocity in km s^{-1} .

On the basis of the figures produced by our algorithm, we classified the 28 stars in the three groups described in Sect. 4. Most of the stars (21) belong to Group 1; we show three example figures of this class below, and the remaining figures and detailed comments can be found in Appendix A. The remaining seven stars are split between Group 2 (4 stars) and Group

3 (3 stars; see also Fig. 1); for these more special cases, we present and discuss all the corresponding figures in the main text.

4. Results of the individual stars

We already showed in Fig. 1 how our sample relates to the whole sample from [de Castro et al. \(2016\)](#) and the location of the stars also as separated into three groups. Group 1 (Sect. 4.1, Figs. 5–7 and A.1–A.18) includes stars that are in agreement with the appropriate AGB models based on the mass of the initial AGB star estimated by [Jorissen et al. \(2019\)](#) and using a metallicity close to that of the Ba star, as described in Sect. 3. Group 2 (Sect. 4.2, Figs. 8–11) includes four stars that show higher abundances of their first s -process peak elements than typically predicted by models close to M_{AGB}^{J19} . Group 3 (Sect. 4.3, Figs. 12–14) includes three stars for which, based on our comparison, M_{AGB}^{J19} may have been overestimated.

4.1. Group 1: Abundance patterns in agreement with M_{AGB}^{J19}

The first group contains stars for which the s -process abundance pattern, and Rb in particular, is in agreement with the AGB mass derived from the binary parameters. We categorised these stars into three sub-groups, depending on the deviations of heavy elements other than Rb. The first sub-group (a) includes stars that have the whole heavy elements pattern in agreement with the models. Stars showing higher Nb and sometimes Mo and Ru abundances than the model predictions belong to the second sub-group (b), while the third sub-group (c) shows abundances

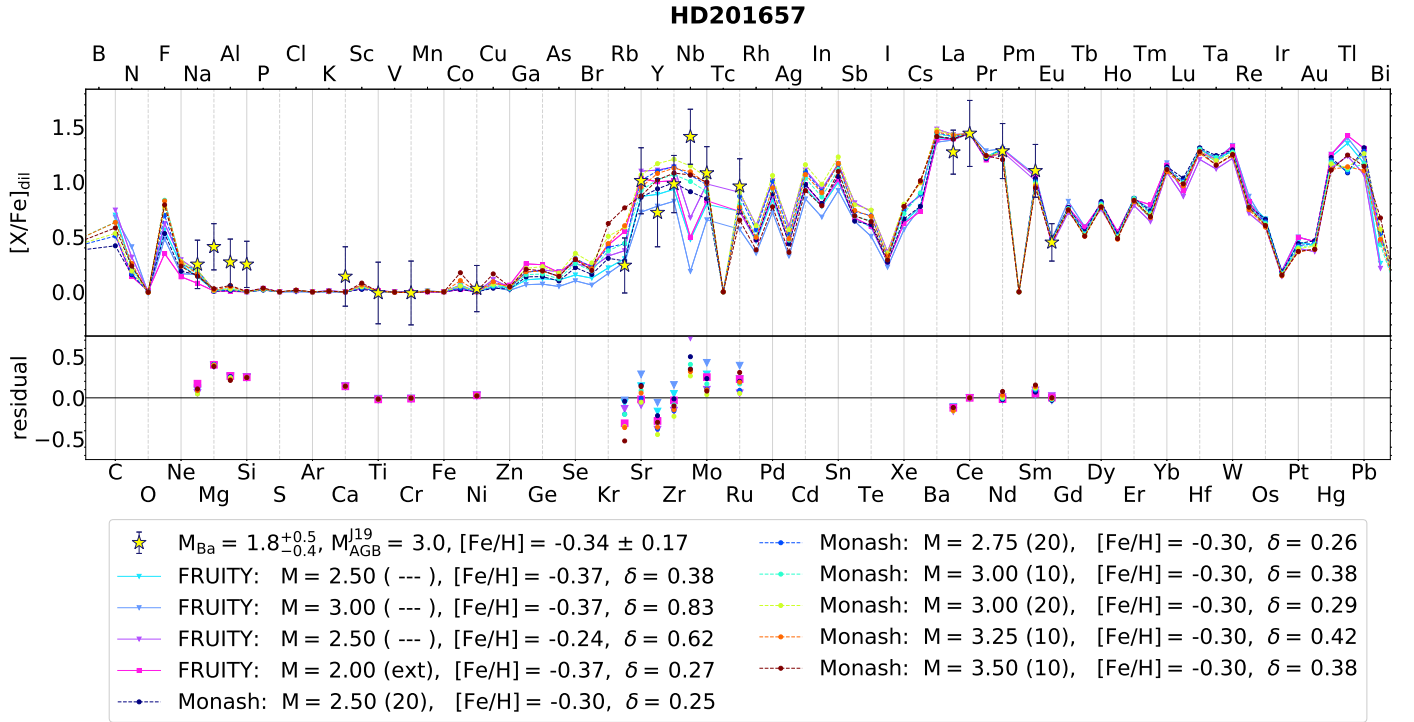


Fig. 6. Same as Fig. 5, but for HD 201657.

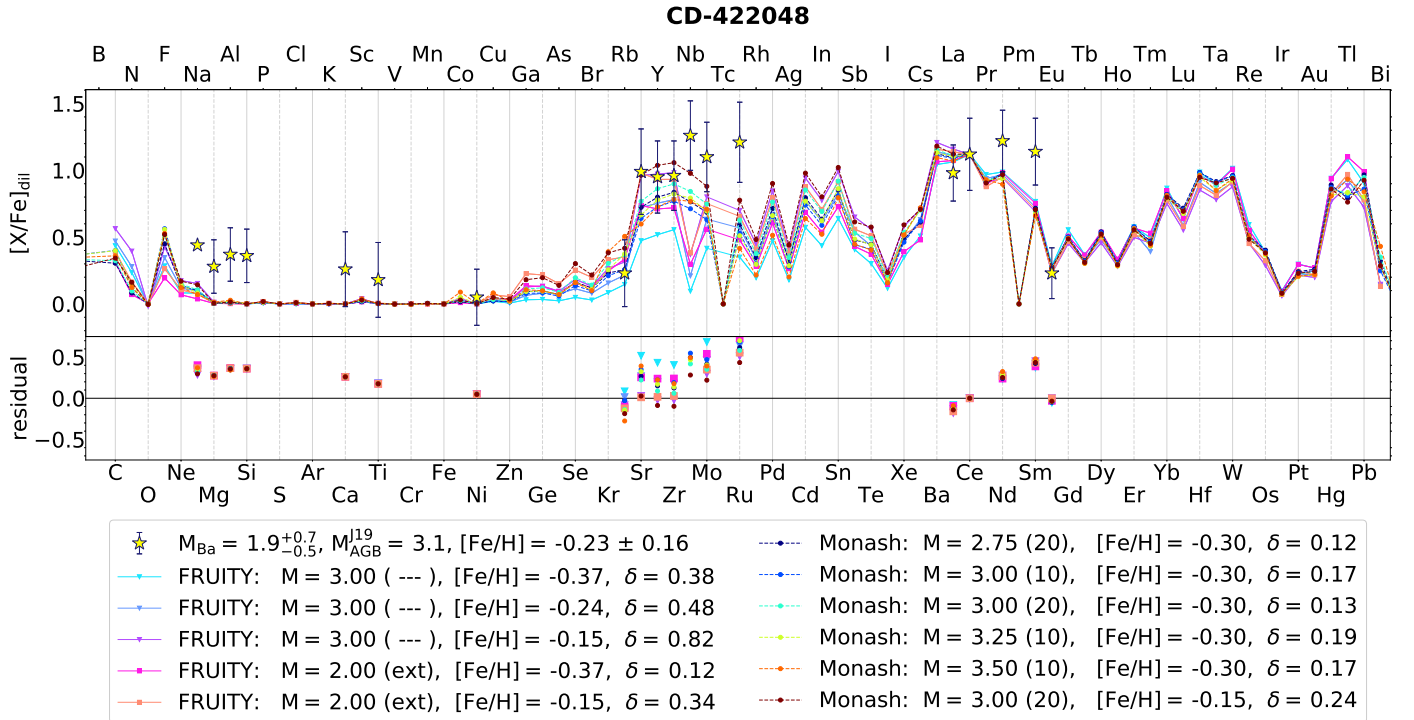


Fig. 7. Same as Fig. 5, but for CD -42°2048.

higher than the model predictions in the elements Nb, Mo, Ru, Nd, and Sm. Below, we discuss one example of each sub-group, while the rest of the sample stars for this group are shown in Appendix A. We refer the reader to Sect. 5 for a general discussion of the *s*-process abundance pattern seen in the sample stars compared to the models, along with the discussion of the light elements and the δ values.

Concerning sub-group (a), the Ba stars HD 58368, HD 91208, HD 154430, and HD 201824 (Figs. 5 and A.1–A.3) belong to this subgroup. As an example, HD 154430 (Fig. 5) is a representative of the sub-group for which most of the elements agree well with the diluted model pattern. Models between 2.0 and 3.75 $M_{\odot}$ are shown. $[La/Fe]$ is somewhat lower than the model predictions, but all the other elements

belonging to the s -process peaks are in agreement with the models. Although $M_{\text{AGB}}^{\text{J19}}$ is $3.6 M_{\odot}$, the large uncertainty of the Ba star mass and the [Rb/Fe] value indicates a polluting AGB with initial mass of $2.5\text{--}3.0 M_{\odot}$. We note that FRUITY models show much larger δ values than the Monash models for this star.

12 stars (HD 20394, HD 40430, HD 53199, HD 58121, HD 59852, HD 95193, HD 119185, HD 143899, HD 200063, HD 201657, HD 210946, and HD 211594, Figs. 6 and A.4–A.14) belong to the sub-group (b) with higher Nb compared to the models, and HD 201657 (Fig. 6) shows an example of this case. We show different models between 2.0 (for the FRUITY ‘TAIL’ case) and $3.5 M_{\odot}$. The abundance pattern (both s -process peaks and Rb) points to a $2.5\text{--}3.0 M_{\odot}$ initial AGB, in agreement with $M_{\text{AGB}}^{\text{J19}}$.

We found that five stars belong to subgroup (c) (BD $-14^{\circ}2678$, CD $-42^{\circ}2048$, CPD $-64^{\circ}4333$, HD 24035, and HD 183915; Figs. 7 and A.15–A.18). They show overabundance in the elements Nb, Mo, Ru, Nd, and Sm compared to the models, and CD $-42^{\circ}2048$ (Fig. 7) is selected as the example. Models in the range close to $M_{\text{AGB}}^{\text{J19}}$ agree well with the observations, although lower masses are also allowed when considering FRUITY models with extended pocket size. We also considered these models because there is a large error bar in the Ba star mass, which would translate into a higher uncertainty in the mass of the initial AGB as well. The $3 M_{\odot}$ FRUITY models show much larger δ values than the Monash models for this star. The [Rb/Fe] value indicates that the mass of the polluting AGB was not above $3 M_{\odot}$. The $3 M_{\odot}$ Monash models with larger ^{13}C pocket have a higher first s -process peak, and thus match the abundances the best. Both of the s -process peaks are in good agreement with the models; however, the elements listed above and located just after the s -process peaks show somewhat higher abundances than the model predictions.

4.2. Group 2: Higher first s -process peak than predicted by the models

The second group consists of stars for which the first s -process peak elements are typically higher than the model predictions, and specific solutions need to be applied. Interestingly, all of the stars belonging to this group have a low $M_{\text{AGB}}^{\text{J19}}$, $<2 M_{\odot}$. Two of them have the lowest AGB determined mass of the whole sample (1.25 and $1.35 M_{\odot}$), and the third and fourth stars have AGB mass $1.8 M_{\odot}$; but, as discussed below in Sect. 4.2.1, they behave differently from the other four stars with similar AGB masses, which belong instead to Group 1 (Figs. A.10, A.13, A.17, and A.18).

4.2.1. HD 18182 (Fig. 8) and HD 180622 (Fig. 9)

For these stars, we show models between 1.5 and $2.25 M_{\odot}$. $M_{\text{AGB}}^{\text{J19}}$ is estimated to be around $1.8 M_{\odot}$, and the nearly solar metallicity of HD 180622 makes it an ideal target to compare different models and model parameters. The low [Rb/Fe] is in agreement with a low-mass polluter AGB. Both Ba stars have a low [Ce/Y] ratio, with HD 180622 showing the lowest [Ce/Y] ratio in our sample. Monash models with different mass and ^{13}C pocket sizes show very similar abundance-pattern and δ values, but these models do not reach the high first s -process peak. To match HD 18182, only a very specific solution is allowed: that of the $2.0 M_{\odot}$ FRUITY T60. For HD 180622, the r30 and r10 models are also a possible match, and we show the [Fe/H] = -0.15 FRUITY T60 models for comparison; although, this is outside of the metallicity range.

The δ value is higher for the r30 model than for r10, with a more feasible value of $\delta = 0.2$.

We conclude that an AGB around $2.0 M_{\odot}$ with ‘TAIL’ or slow rotation could be the polluter of these Ba stars. These solutions are possible because these models have a lower neutron exposure, which favours the production of the first peak elements relative to the second peak elements. However, for HD 18182 these models are still factor of 3 to 5 lower than the observed [Mo/Fe] and [Ru/Fe] values, as is the case with the stars of sub-group 1b, while HD 180622 also has higher Nd than the model predictions, similar to the stars in sub-group 1c.

4.2.2. HD 107541 (Fig. 10)

The AGB companion of HD 107541 has a very low $M_{\text{AGB}}^{\text{J19}}$ of $1.4 M_{\odot}$. Monash models are presented between 1.5 and $2.0 M_{\odot}$, and we plot two FRUITY $2 M_{\odot}$ model with different pocket sizes. Except for the $2 M_{\odot}$ FRUITY model with a standard pocket size, all the other models can match Rb, Sr, and Zr within the error bars. Only the $2 M_{\odot}$ Monash model with a larger pocket size can match the abundances with a reasonable δ value, but this mass is far away from $M_{\text{AGB}}^{\text{J19}}$. Mo and Ru are overabundant compared to the models, which is similar to the case of the stars belonging to sub-group 1b.

4.2.3. HD 134698 (Fig. 11)

The star HD 134698 also has a very low $M_{\text{AGB}}^{\text{J19}}$ of $1.2 M_{\odot}$, and we show models up to $1.75 M_{\odot}$. The [Rb/Fe] value and the first s -process peak shows a higher value than the presented model masses, close to $M_{\text{AGB}}^{\text{J19}}$. The metallicity of the models is slightly lower than that of this Ba star; however, we checked the FRUITY models with [Fe/H] = -0.37 , and we found that they also predict a higher second-to-first peak ratio than observed in this star. A match to the [Rb/Fe] value can be found by higher mass models, although none of these models can reproduce the high abundance values of the other light s -process elements. We also note that this star has the most eccentric orbit ($e = 0.95$, Jorissen et al. 2019) among our sample stars.

4.3. Group 3: Abundance patterns that require lower stellar masses than $M_{\text{AGB}}^{\text{J19}}$

Three of the sample stars appear to have overestimated $M_{\text{AGB}}^{\text{J19}}$ (between 3.8 and $5.6 M_{\odot}$) when comparing the abundance pattern with the models. This is because models of the corresponding masses have typically smaller ^{13}C pockets than their lower masses counterparts (see discussion in Cristallo et al. 2009; Karakas & Lugaro 2016) and therefore do not produce enough s -process element abundances to reach the very high observed values of, for example, [Ce/Fe] between 1 and 2 dex. Below we show models for each star that match the abundance pattern. Our suggested initial AGB masses are between 2 and $3 M_{\odot}$ for each star in this group. Furthermore, Rb is available for two of these stars and is lower than the higher mass model predictions.

4.3.1. HD 49641 (Fig. 12)

The star HD 49641 is the first star in our sample that has a high $M_{\text{AGB}}^{\text{J19}}$ of $5.2 M_{\odot}$. Considering the large error bar on the mass of the Ba star, we searched for matching AGB models with masses down to $2.0 M_{\odot}$. Unfortunately, no [Rb/Fe] values are available for this star to better constrain the initial AGB mass.

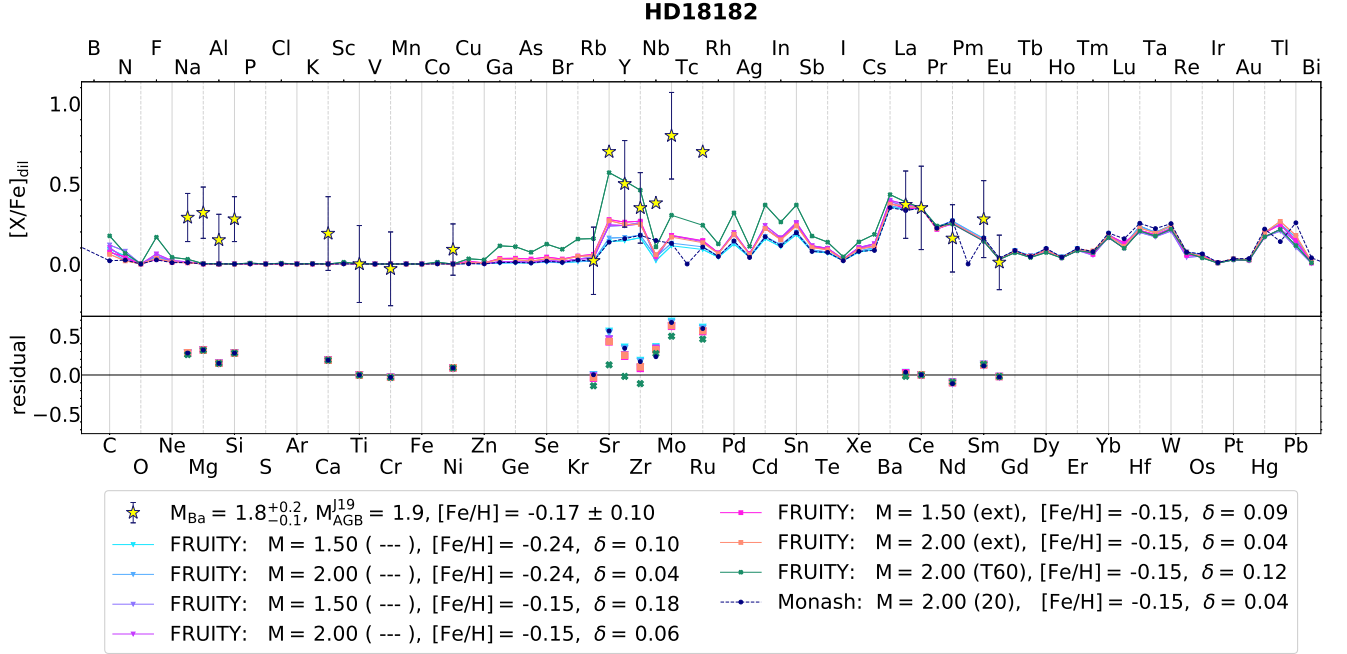


Fig. 8. Same as Fig. 5, but for HD 18182.

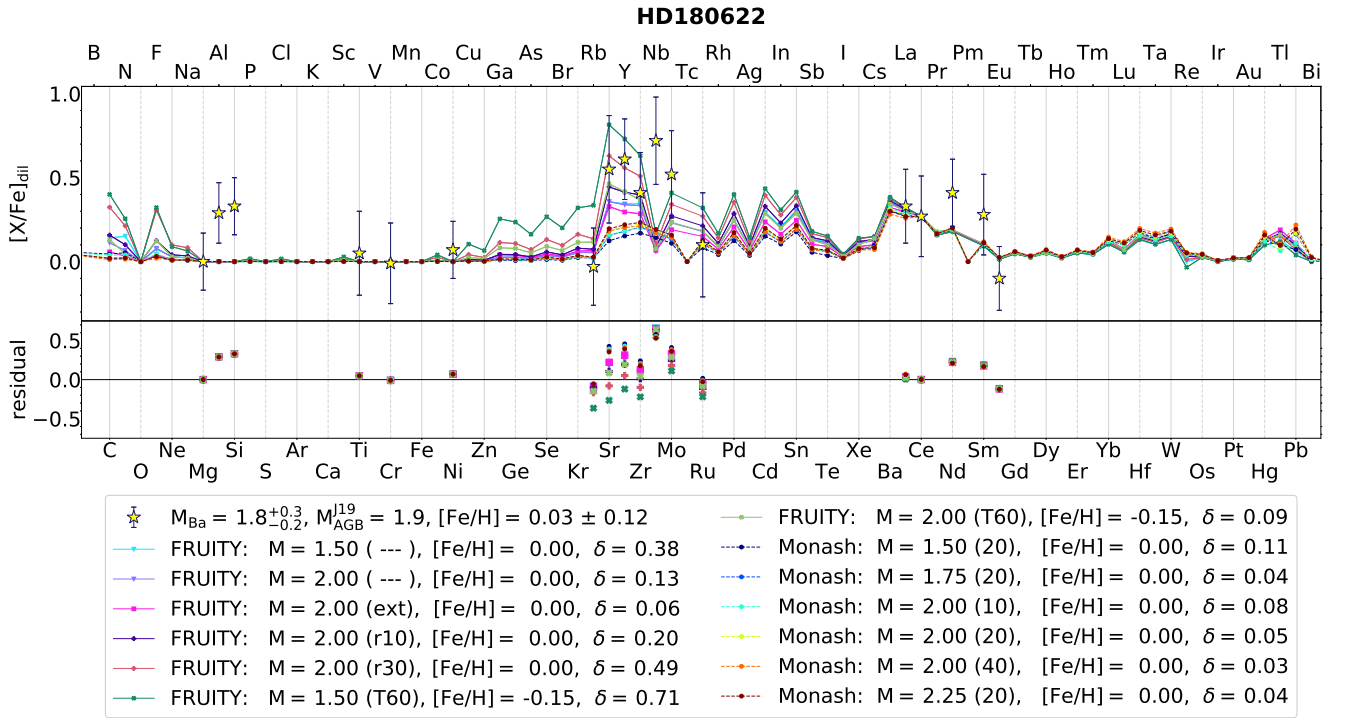


Fig. 9. Same as Fig. 5, but for HD 180622.

Additionally, the error bars of the first *s*-process peak elements are covering most of the models shown here, which does not allow us to constrain the mass of the polluting AGB.

4.3.2. HD 84678 (Fig. 13)

We show four different models for HD 84678 between 2.5 and 4.5 $M_{\odot}$. The larger mass models overestimate the $[Rb/Fe]$ value, which points to a much lower initial AGB mass, below 2.5 $M_{\odot}$. Only the 2.5 $M_{\odot}$ FRUITY model is close to the abundance pat-

tern, although it has a very high δ . Furthermore, none of the models match the overabundance of Nd and Sm, as is the case for stars in sub-group 1b.

4.3.3. HD 92626 (Fig. 14)

The star HD 92626 has an M_{AGB}^{19} of 5.6 $M_{\odot}$. We show models between 2.0 and 3.0 $M_{\odot}$; since the $[Rb/Fe]$ value is pointing to such low mass model, the 3 $M_{\odot}$ models already show much higher Rb values than observed. We note that none of the

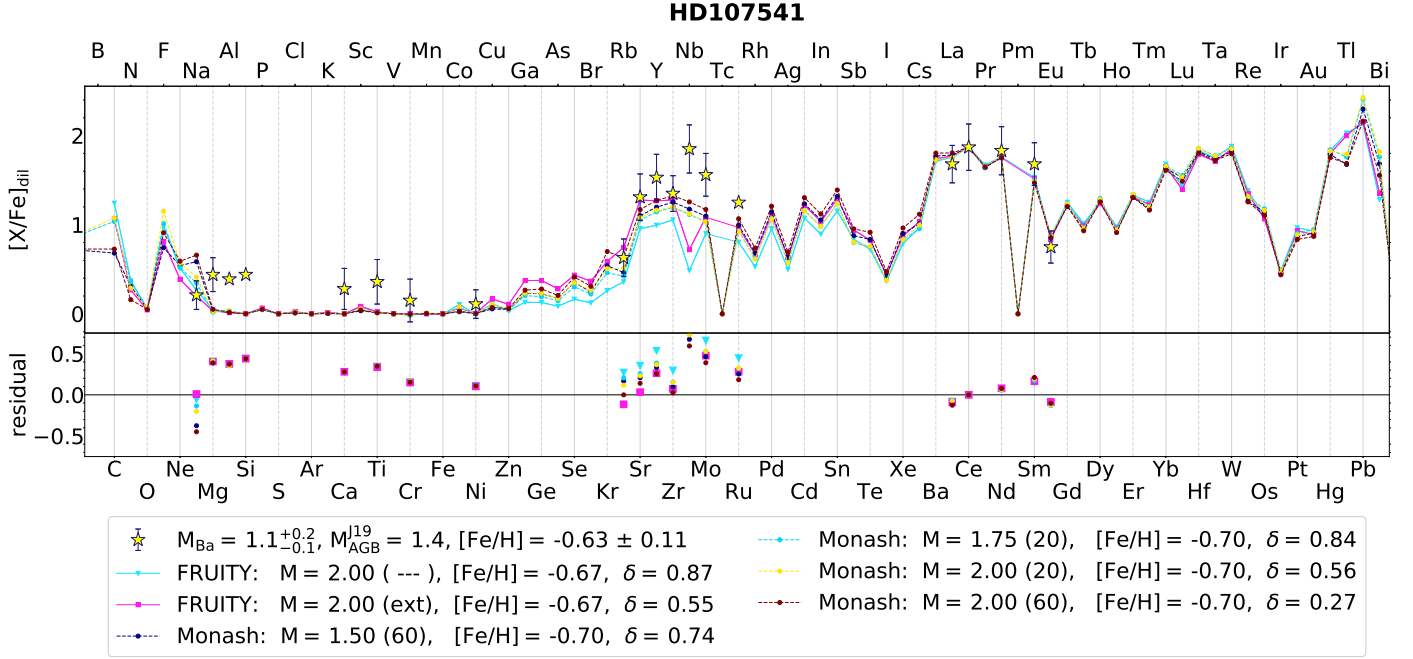


Fig. 10. Same as Fig. 5, but for HD 107541.

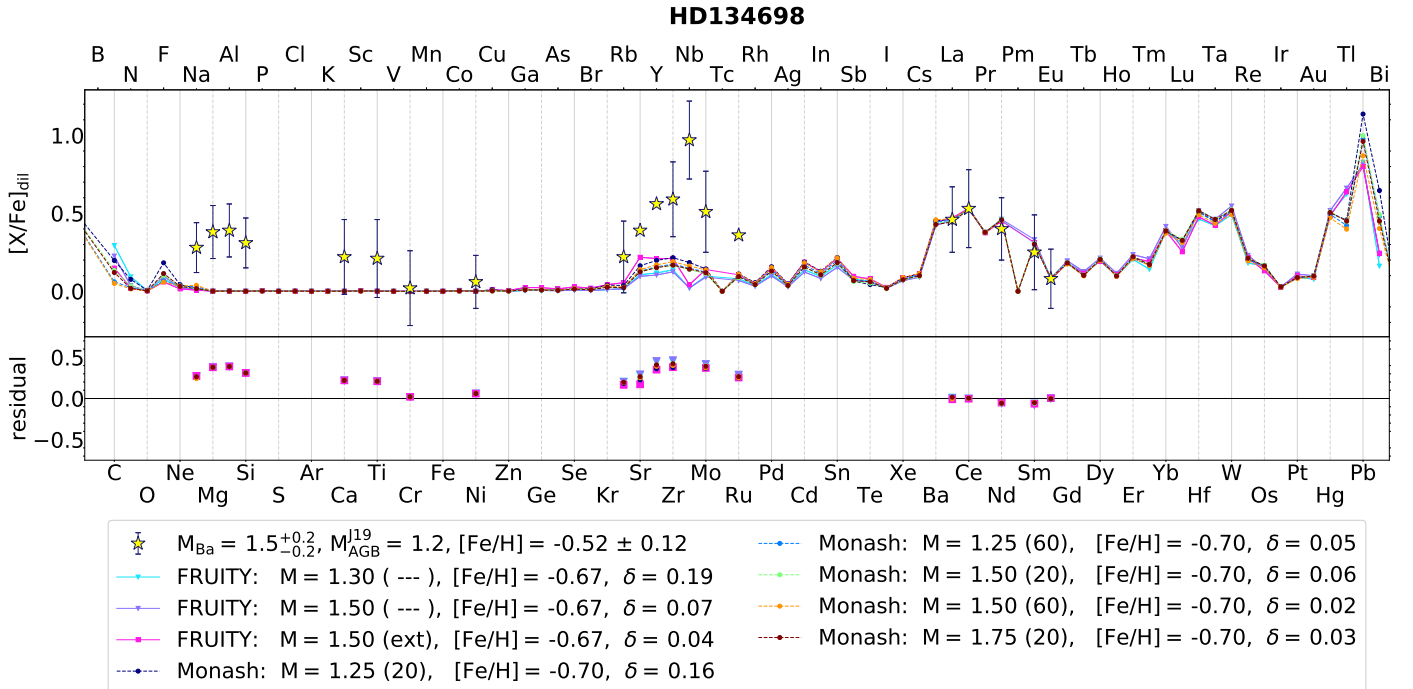


Fig. 11. Same as Fig. 5, but for HD 134698.

models with metallicity higher than those shown in the figure can match the observed pattern within our dilution limit ($\delta \leq 0.9$). This indicates that the mass of the polluting AGB may be overestimated and was between $2\text{--}2.5 M_{\odot}$, and with metallicity of ≈ -0.30 dex, which is within the given error bar for this star. To match the observed abundance pattern, high δ values, and thus a large amount of accreted AGB mass, are needed. The orbital period of this system being relatively short (918 days) could explain the higher amount of AGB material in the atmosphere of the Ba star, as compared to other, longer period ($\approx 5\text{--}10$ yr) systems.

5. Discussion

As discussed above, the $[Rb/Fe]$ value of the individual stars is a good proxy for the initial mass of the polluter AGB (Roriz et al. 2021a). The $[Rb/Fe]$ of the star generally does confirm M_{AGB}^{19} . Fifteen stars in our sample were also included in the work of Stancliffe (2021), who considered the Ba stars in the sample of de Castro et al. (2016) with $[Fe/H] = -0.25 \pm 0.125$ dex and, using the Monash models, found that most of these stars can be reproduced by AGB companions of 2.5 and $3 M_{\odot}$ in mass, similar to what we report here.

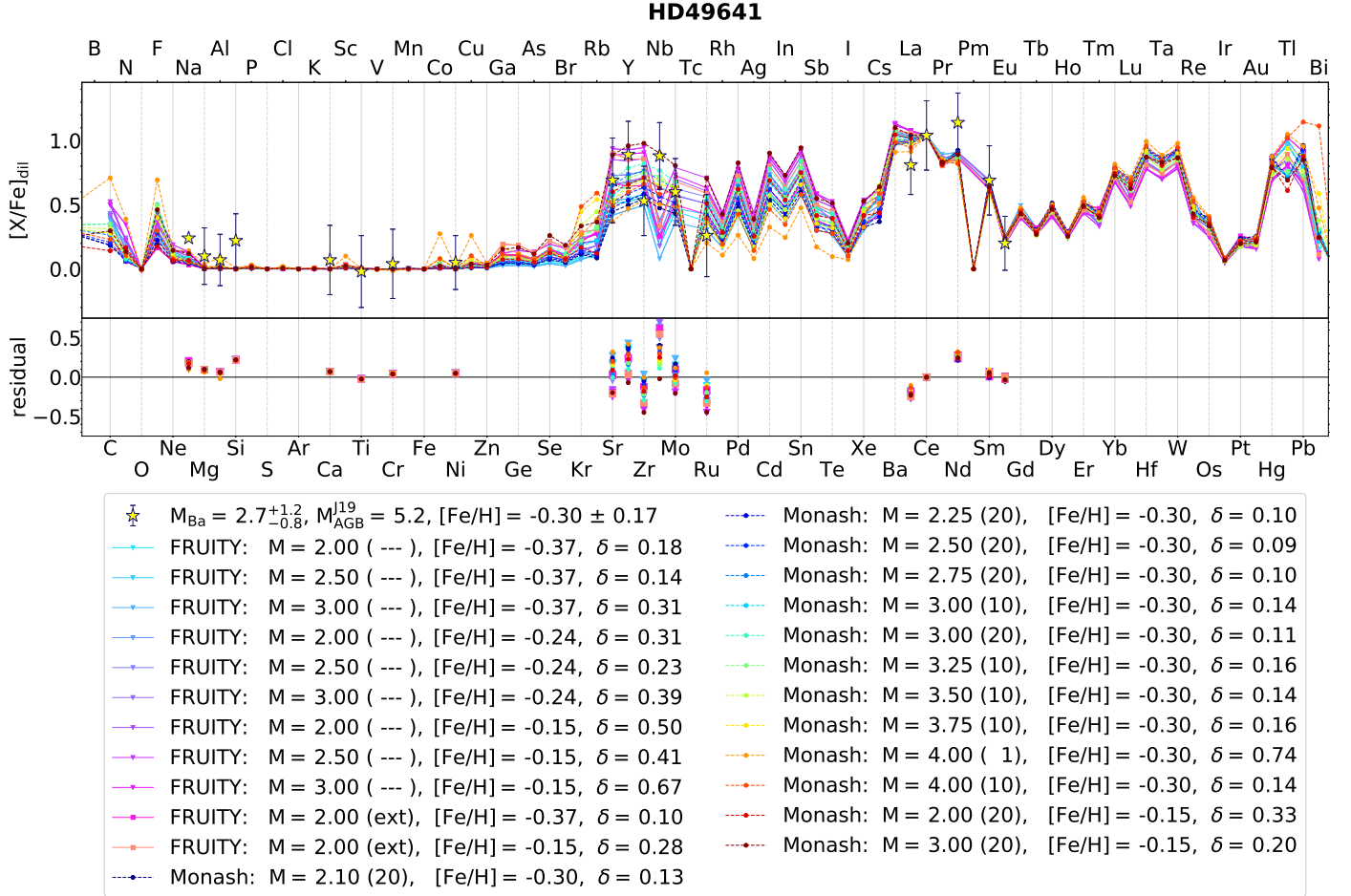


Fig. 12. Same as Fig. 5, but for HD 49641.

Among the Group 3 stars specifically (see Sect. 4.3), there are instead discrepancies between the initial masses derived by Jorissen et al. (2019) using the constant Q method and the AGB mass indicated by the $[Rb/Fe]$ value. When applying the constant Q assumption for the determination of the WD masses, some stars have WD masses close to or above $1 M_{\odot}$ (initial AGB mass $\approx 5.7 M_{\odot}$). However, we did not expect for Ba stars to be enriched by such a high mass companion at metallicity close to solar. We found that the polluter AGB might also have been between 2 and $3 M_{\odot}$ for these stars, instead of a more massive star.

The shape of the first s -process peak elements (Sr, Y, Zr) shows different pattern for each star; for example, in some cases the Zr abundance is much lower than the neighbouring Y values. We should note that the Zr abundances used in this study were derived from Zr I lines, which are systematically lower than those derived from Zr II lines. Such a fact is generally attributed to departures from the LTE approximation, as discussed by certain authors (e.g. Gratton & Sneden 1994; Allen & Barbuy 2006; Karinkuzhi et al. 2018). Further investigation aiming to compare abundances from Y I and Zr II lines for the sample of de Castro et al. (2016) could be useful in order to quantify these differences.

The pattern of the second peak elements (La, Ce, Nd) is in good agreement with the models shown here. This is not surprising, since we normalised the models to match the $[Ce/Fe]$ values of the sample stars. However, for some stars the $[La/Fe]$ value is somewhat lower than model predictions. Since we did not take into account the uncertainties of $[Ce/Fe]$, lowering

the models within the uncertainties of $[Ce/Fe]$ could solve this issue.

Eu is mostly produced by the r process, and the $[Eu/Fe]$ abundances are in most cases in agreement with the model predictions. The overall distribution of $[Eu/Fe]$ abundances of the sample stars is similar to the disc population (Roriz et al. 2021b).

In relation to Group 2, Shetye et al. (2019) recently found S-type stars with initial masses of around $1 M_{\odot}$ in the metallicity range between -0.27 and -0.54 dex and enhancements in s -process elements requiring the TDU to also operate at these masses. Therefore, AGB stars of around $1 M_{\odot}$ may be able to produce heavy elements at the beginning of the AGB phase and the polluter AGB stars of the Ba stars belonging to our Group 2 could represent other examples of this group of low mass AGBs with active TDU. As can be seen in Fig. 11 for HD 134698, unlike the models of Shetye et al. (2019) the low-mass models used in this study do not reach the large enhancement in the first s -process peak elements. In fact, the dredge-up at these lower masses is very model-dependent: higher metallicity Monash models below $1.5 M_{\odot}$ do not have a PMZ leading to the formation of s -process elements at given masses, while FRUITY $1.3 M_{\odot}$ models only have high enough enrichment in s -process elements ($[s/Fe] \geq 0.25$) for metallicities below $[Fe/H] = -0.24$.

When examining the orbital parameters of our sample stars (Fig. 15), we see that the three stars in Group 3 have eccentricities below 0.1 and periods below 2000 days. While there are three stars in Group 1 in the same region of Group 3, all stars in Group 2 instead have periods longer than 3000 days. HD 134698,

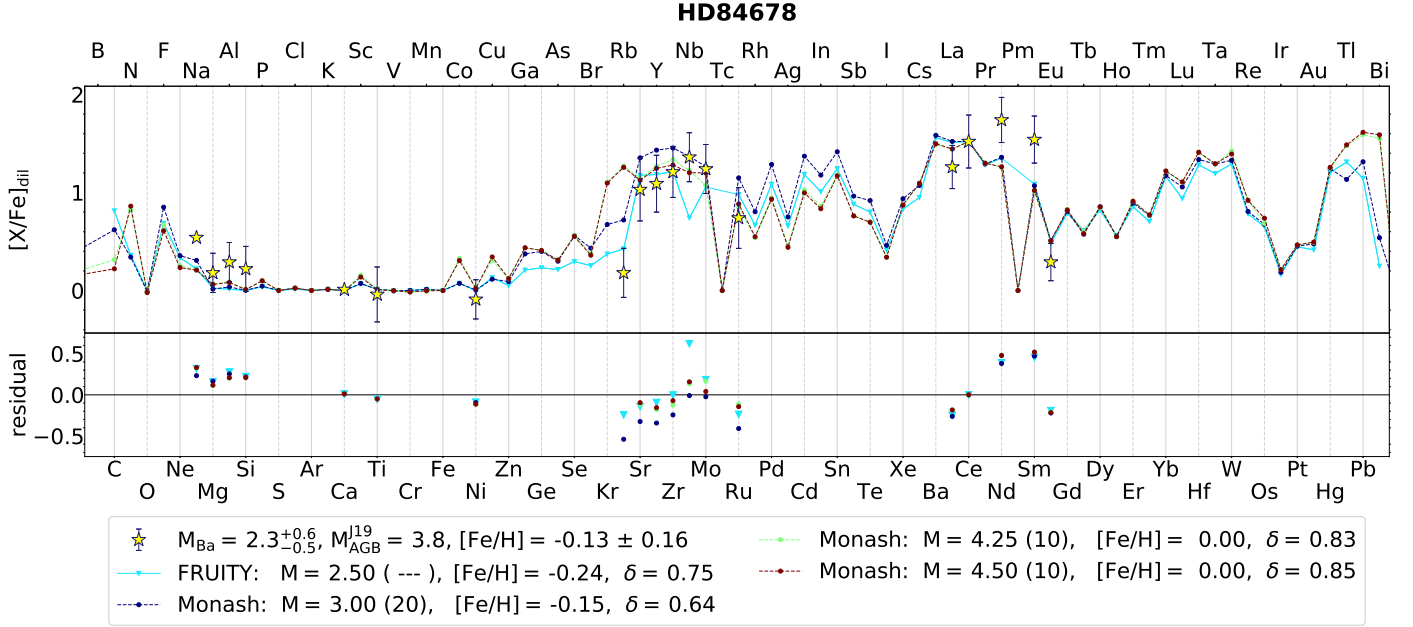


Fig. 13. Same as Fig. 5, but for HD 84678.

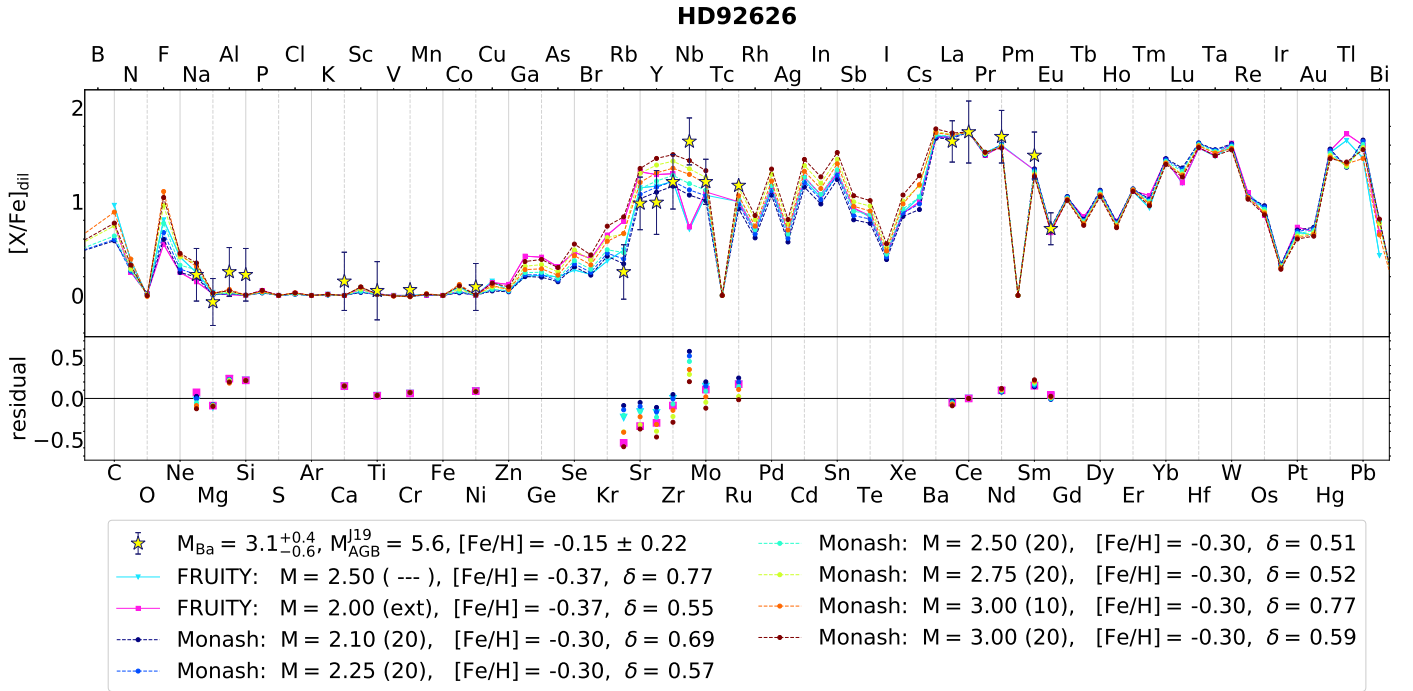


Fig. 14. Same as Fig. 5, but for HD 92626.

the most eccentric system in our sample, also belongs to Group 2. The other three stars in Group 2 share similar binary parameters to some stars in Group 1. We have to note that some of our sample stars only have preliminary orbital parameters (see Table 4 in [Jorissen et al. 2019](#)).

5.1. The problem of sub-groups 1b and 1c

The stars belonging to sub-groups 1b and 1c, as well as the stars of Group 2, show excess in the heavy elements that immediately follow the first and second s -process peaks compared to the models. As already discussed in [Roriz et al. \(2021b\)](#), their Fig. 9) the

explanation of the excess of Nb, Mo, and Ru relative to Sr and of Nd and Sm relative to Ce, within the framework of the spectroscopic analysis, is not trivial.

The fact that these elements are produced not only via the s process predicts that contributions from other stellar sources or nucleosynthetic processes might be missing from the AGB models used in this analysis. For example, some stars show an $[Sm/Fe]$ value higher than the model predictions. Since Sm is produced by $\approx 70\%$ the r process, the slight overabundance can be caused by a missing r -process component of this element in the AGB models. Nuclear uncertainties could also play a role in the deviations, although neutron capture cross sections are

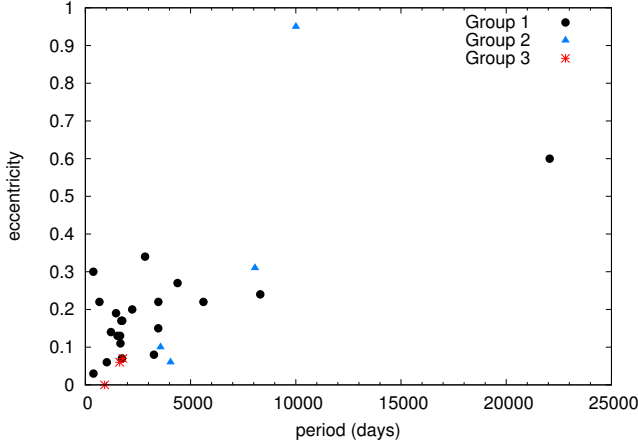


Fig. 15. Period-eccentricity distribution of the sample stars.

relatively well known for stable isotopes. However, as pointed out by [Roriz et al. \(2021b\)](#), the overabundance of Sm is so high that there are probably no signs of the initial abundances left in the observed abundance pattern.

The study of [Mishenina et al. \(2019\)](#) showed that the available galactic chemical evolution models do not produce enough Mo and Ru compared to the disc star observations in the metallicity range of our sample stars. In the galactic chemical evolution models of [Kobayashi et al. \(2020\)](#), Mo and Ru are both overproduced compared to the observations when specific components are added to the model computations. These uncertainties still leave the question of the observed overabundance of certain elements open.

An explanation for these stars might be the activation of the intermediate neutron capture process (*i* process), which was introduced to explain several peculiarities observed in different types of stellar objects (e.g. CEMP stars, post-AGBs, and also the extreme Ba enhancement in young stars [Baratella et al. 2021](#)). The *i* process has higher neutron densities compared to the *s*-process, which typically manifests in $[\text{Rb}/\text{Sr}] \geq 1.0$. However, this is not the case for our sample Ba stars, all stars having negative $[\text{Rb}/\text{Sr}]$. Late neutron fluxes occurring in some AGBs might be able to keep Rb low compared to Sr and produce more proton-rich elements close to the first and second peaks by shifting part of the elements to heavier ones, from Sr, Y, Zr to Nb, Mo, Ru and from Ba, La, Ce to Nd, Sm, for the first and second peaks, respectively. The exact process needs further investigation.

5.2. The light elements and C/O ratio

Although C and O were not derived for the sample stars, we briefly discuss these elements to stress their importance in a future analysis. As is clear in the figures, some of the stars require high δ values to be matched, and this may create problems with respect to the C/O ratio. If $[\text{C}/\text{Fe}] > 0.27$, this results in Ba stars that are C rich ($\text{C} > \text{O}$); however, these stars are not observed to be C-rich. For example, if we wish to match HD 91208 (Fig. A.2) using the FRUITY $2.5 M_{\odot}$ model, we obtain $\delta = 0.67$. However, such a high value of δ results in $[\text{C}/\text{Fe}] \approx 0.5$, which corresponds to $\text{C}/\text{O} \approx 2$ and would lead to this Ba star being observed as a C-star instead. Therefore, such large δ are not realistic from the abundances point of view either. This problem could be mitigated in some cases if the initial O abundance, which is not observed and is unaffected by AGB

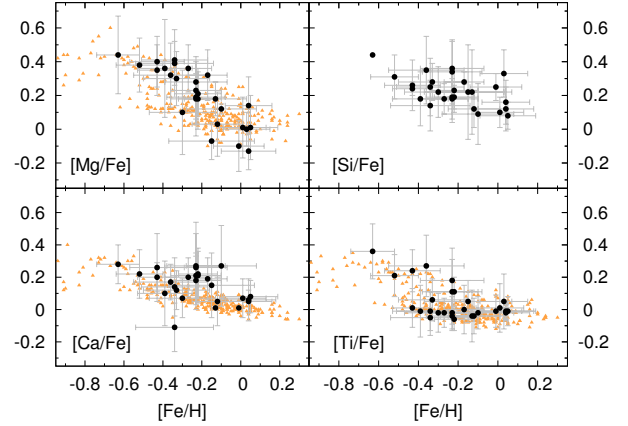


Fig. 16. α -element abundances for the sample Ba stars (black dots) and literature data for solar neighbourhood giants (orange triangles) from [Jönsson et al. \(2017\)](#).

stars, is higher than solar. This would keep the C/O ratio lower, and it applies mostly to stars of metallicity lower than solar, due to the effect of α -enhancements. This effect can be derived from the observed abundances of the other α -elements, especially Mg, Si, Ca, and Ti, which may be higher compared to AGB models. These elements reflect the initial composition of the stars and vary from star to star due to their different birthplaces. This was already pointed out by [de Castro et al. \(2016\)](#), and they suggested that some stars are in transition between the thin and the thick disc. For example, HD 201657 (Fig. 6) has $[\text{Mg}/\text{Fe}] \approx 0.4$, and if we also assume $[\text{O}/\text{Fe}] \approx 0.4$, the $2.5 M_{\odot}$ FRUITY models with relatively high δ values and high $[\text{C}/\text{Fe}]$ could still produce a solution with an O-rich star. Instead, in relation to the above example of HD 91208, this star has solar metallicity and these elements are not enhanced relatively to solar.

Since the sample stars are all located in the solar neighbourhood, we can also compare the distribution of their light elements to that of normal giants. We found that the Ba stars examined in this study show similar α -element abundances to the normal giants in the solar neighbourhood analysed by [Jönsson et al. \(2017\)](#). Figure 16 shows that the abundances of these elements in the Ba stars increase when the metallicity decreases, in agreement with solar neighbourhood giants and galactic chemical evolution models.

5.3. The δ values

Figure 17 shows the average δ values over all the matching models versus period for the sample stars. The error bar on the δ value is the standard deviation, and some of the plotted FRUITY models were excluded from the calculation. The T60 models were not included for HD 40430, HD 53199, HD 59852, HD 143899, HD 180622, and HD 200063, since in these cases this model is only plotted for comparison and completeness, but it is not a good match to the abundance pattern. The removal of the FRUITY r30 and $[\text{Fe}/\text{H}] = 0.15$ dex models for HD 58121 and HD 95193, respectively, lowers the average δ value below the 0.3 limit for these two stars. Overall, the average δ value is lower than ≈ 0.3 for all the systems with period over ≈ 4000 days, while stars with short periods (< 1000 days) have $\delta \geq 0.3$, although we should note that some longer period systems have uncertain periods.

Overall, though depending on the exact parameters, FRUITY models need higher δ values to match the observed abundances

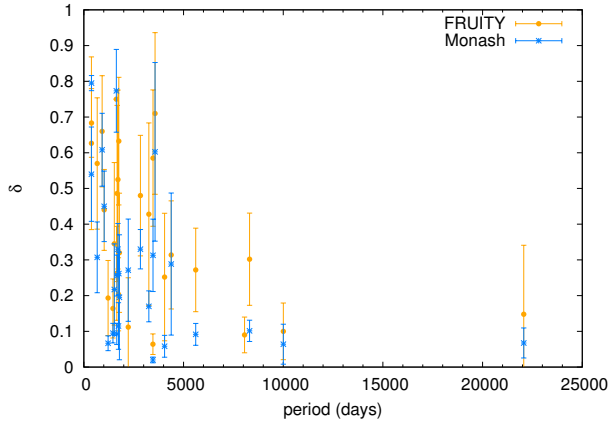


Fig. 17. δ values as a function of the orbital period separately for FRUITY and Monash models. The error bar on the δ value is the standard deviation.

pattern. This is because, as shown in Fig. 4 and mentioned in Sect. 3.2, these models predict typically lower abundances at the AGB surface. In the Monash models using a larger ^{13}C pocket of course leads to a reduction of the δ value, since there is more Ce produced.

Most of the stars with binary period in the typical range of 1000 to 10 000 days can be fitted by models with δ values below 0.2, except for HD 211594, HD 201824, and HD 201657 (all belonging to Group 1), which need δ values from 0.3 and above. Instead, the three stars (CPD $-64^\circ 4333$, HD 24035 and HD 92626) with the shortest orbital periods (of less than 3 yr) typically require models with δ values higher than roughly 0.4. For example, for CPD $-64^\circ 4333$, all of the matching models have δ values above 0.6. Together with HD 24035, CPD $-64^\circ 4333$ belongs to Group 1c, while the third star, HD 92626, belongs instead to Group 3. HD 84678, belonging to Group 3, also needs δ values above 0.6 to be matched, even though its period is 1630 days. Among the Group 2 stars, which have the lowest masses, only HD 107541 requires relatively high δ values above 0.55. It would be interesting to further analyse the binary evolution of all these peculiar systems individually. The fact that these stars need higher δ values may be related to a different mass transfer and binary evolution, rather than the AGB model, as also discussed by Stancliffe (2021) for HD 24035. For example, based on hydrodynamical simulations of binary AGB systems, Chen et al. (2020) found that with decreasing binary separation, the accretion rate increases. In their simulations, the two least separated binaries have dusty circumbinary discs, and these authors showed that if a circumbinary disc is present in the system, the mass transfer efficiency (defined as the ratio of the accretion rate of the secondary, the current Ba star, and the mass-loss rate of the AGB primary) through wind-Roche-lobe overflow can increase up to 31%, eight times larger than the canonical Bondi-Hoyle-Lyttleton accretion mechanism. Thus, the presence of a circumbinary disc during the mass transfer may explain the higher δ values for the systems, where even the largest Monash ^{13}C -pocket size model does not give a satisfactory match.

6. Conclusions

We analysed the individual abundance patterns of 28 Ba stars against AGB model predictions of the *s* process. For these stars, masses were independently derived by Jorissen et al. (2019) via

an investigation of binary interaction, and high-resolution spectroscopy abundances were reported by de Castro et al. (2016). New elemental abundances for Rb (Roriz et al. 2021a), Sr, Nb, Mo, Ru, La, Sm, and Eu (Roriz et al. 2021b) allowed us to carry out a comparative study between the individual abundance patterns with different AGB models in the metallicity range of the stars using dilution to match the [Ce/Fe] values of each star.

We grouped our sample stars into three main groups. The first group includes 21 stars that have matching AGB models, the second group includes four stars with higher first *s*-process peak abundances compared to most models, and the third group includes three stars with high M_{AGB}^{119} ($>3.8 M_\odot$), while our investigation supports a lower mass. Our main conclusions are the following:

1. Stars in Group 1 are in good agreement with the models, but stars in sub-groups b and c show overabundances in Nb and in Nb, Mo, Ru, Nd, and Sm, respectively. This overabundance was already pointed out in the large sample analysed by Roriz et al. (2021b), who suggested a different nucleosynthesis path to reproduce the excess of these heavy elements located just after the *s*-process peaks. Our analysis of individual stars confirms the need for such future investigation.
2. Group 2 includes four stars with low estimated Ba and AGB mass ($<2 M_\odot$) and with higher first *s*-process peak than typically predicted, which requires somewhat lower neutron exposures, as found in the FRUITY rotating or ‘TAIL’ models. However, there are only two stars out of the four in Group 2, where these models can give a possible solution to match the observed abundance pattern. The production of the light *s*-process elements can also be attributed to diffusive mixing (Battino et al. 2019), and these models will need to be included in a future study.
3. Group 3 includes three stars in our sample that have higher initial AGB mass ($>3.8 M_\odot$) derived under the constant Q assumption (Jorissen et al. 2019). As these authors also indicated, these AGB masses might be somewhat overestimated due to the high WD mass. For these stars, we found that lower mass AGB models (below $4 M_\odot$) can explain the elemental patterns.
4. Most of the light elements in our sample stars show excess compared to the AGB models, but they are in agreement with the overall trend with the solar neighbourhood giants and with the predicted trend of galactic chemical evolution.
5. The calculated δ values are mostly lower than 0.3. However, some systems need larger δ values, and thus probably higher mass transfer efficiency to reproduce the Ba star abundance pattern. These stars also have typically shorter orbital periods, which might explain the larger amount of transferred AGB material.

In this paper, we confirm the findings of Roriz et al. (2021a,b) using the abundance pattern of the individual stars rather than comparing the bulk sample of the Ba stars to AGB models. None of the stars show high Rb or $[\text{Rb}/\text{Sr}] > 0$ dex (Roriz et al. 2021a), which means that AGBs below $4 M_\odot$ are the contributors to the abundance pattern of that observed in this Ba star sample. This raises the question of where the Ba stars polluted by a higher mass AGB companions are. To resolve this problem, we should consider including stars with $[\text{s}/\text{Fe}] < 0.25$ dex in the Ba star category, which is expected from the theoretical models of mass over $\approx 4 M_\odot$.

Future work involves further spectroscopic analysis of the stars belonging to Groups 1b and 1c and Group 2 and the comparison with other sets of models. Furthermore, elemental abundances of C and the third peak element Pb could help to

distinguish between the processes in the AGB stars that lead to the formation of the observed Ba stars. Moreover, the dilution due to mass transfer and its implication in relation to each individual binary system should be further investigated by connecting the δ values reported here for each star to a value of the accreted mass and to a possible mass transfer efficiency, which are the quantities used in the modelling of the binary interaction. For example, Stancliffe (2021) presented models of the accretion of AGB ejecta onto low-mass companions to form Ba stars of typical mass ($2.5 M_{\odot}$). Such models could be targeted at each individual system considered here, as the now available and more extended abundance pattern, together with the known binary properties and masses, could help to derive more precise accreted masses. The corresponding binary-derived dilution factors can then be compared to those derived here on the basis of AGB nucleosynthesis. The method used by Abate et al. (2015) for CEMP-s stars could also be applied to the sample of Ba stars with known orbital periods. Finally, to be able to analyse each individual Ba star of the full sample of 169, we are implementing machine learning techniques, which will be presented in a second paper in this series.

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Appendix A: Figures

We present here the results for the rest of the stars belonging to Group 1. The stars are grouped in the same way as in the main text: stars with the whole heavy elements pattern in agreement with the models belong to sub-group 1a, sub-group 1b includes stars with higher Nb and sometimes Mo and Ru compared to the models, and stars in sub-group 1c show abundances higher than the model predictions in the elements Nb, Mo, Ru, Nd, and Sm. The δ values were calculated as described in Sect. 3.2, normalising the models to match the [Ce/Fe] of the individual stars. We only show models with $\delta \leq 0.9$ and with a good match to the observed [Rb/Fe] values.

A.1. Group 1a: Observed abundance pattern in agreement with model predictions

A.1.1. HD 58368 (Fig. A.1)

We show models between 2.0 and 3.5 $M_{\odot}$. These models match the Rb abundance and the second *s*-process peak well, although the observed Sr and Zr abundances are lower than all of the model values. The Rb value indicates that the polluting AGB was around 3 $M_{\odot}$, in agreement with M_{AGB}^{J19} of 3.1 $M_{\odot}$. Higher mass FRUITY models have higher δ values than our limit of 0.9.

A.1.2. HD 91208 (Fig. A.2)

For HD 91208 M_{AGB}^{J19} is close to 3.0 $M_{\odot}$, and we show models between 2.0 $M_{\odot}$ (FRUITY ‘TAIL’ and T60) and 3.25 $M_{\odot}$, which we selected as the best compromise to cover both Rb and the first peak elements Sr and Y. Unfortunately, there is no available error bar for Nb, but its proximity to the models allows us to categorise this star in sub-group 1a. We note that FRUITY models show much larger δ values than the Monash models for this star.

A.1.3. HD 201824 (Fig. A.3)

We show FRUITY models between 2.0 and 2.5 $M_{\odot}$ and Monash models between 2.1 and 3.0 $M_{\odot}$. FRUITY models higher than 2.5 $M_{\odot}$ have larger δ values than our limit of 0.9. The FRUITY 2.5 $M_{\odot}$ model with [Fe/H] = −0.37 and Monash models below 3.0 $M_{\odot}$ match the abundance pattern very well; thus, we conclude that the polluter AGB was around 2.5 $M_{\odot}$, close to M_{AGB}^{J19} .

A.2. Group 1b: Higher observed Nb (and sometimes Mo and Ru) than model predictions

A.2.1. HD 20394 (Fig. A.4)

Both *s*-process peaks are matched well by the models in the mass range of 2–3 $M_{\odot}$. We note that FRUITY models show much larger δ values than the Monash models for this star. The match of [Rb/Fe] with the models points to an AGB with initial mass of 3 $M_{\odot}$ or lower rather than higher. Although Nb is matched by the 3 $M_{\odot}$ Monash model, that model is far from matching Rb; thus, we kept this star in sub-group 1b.

A.2.2. HD 40430 (Fig. A.5)

In agreement with M_{AGB}^{J19} of 2.8 $M_{\odot}$, models between 2.0 and 3.25 $M_{\odot}$ match the measured elements of HD 40430 well. We also show the FRUITY T60 model, which is far from the abundance pattern of the first peak of this Ba star. The Rb value points

to initial AGB masses in the lower mass range, $\approx 2.5 M_{\odot}$. Nb is higher than the model predictions; however, since an error bar is not available, we kept this star in sub-group 1b.

A.2.3. HD 53199 (Fig. A.6)

Models between 2.0 and 3.5 $M_{\odot}$ are shown for HD 53199. Except from the 2.0 $M_{\odot}$ FRUITY T60 model, all models match the *s*-process peaks well, with the lower than predicted [Rb/Fe] value indicating the lowest initial AGB mass. Models above 3.5 $M_{\odot}$ cannot reach the observed [Rb/Fe] and first *s*-process abundances at the same time; thus, we confirm that the polluter AGB is around 3 $M_{\odot}$, in agreement with M_{AGB}^{J19} .

A.2.4. HD 58121 (Fig. A.7)

We show different models with masses between 2.0 and 3.5 $M_{\odot}$. Except for the FRUITY r30 model, which shows higher light *s*-elements than other models, all the remaining models are in good agreement with the *s*-process peak elements, although [Rb/Fe] and [Eu/Fe] are lower than 0 and impossible to match by model predictions. Yang et al. (2016) determined Eu abundance for this star using only one line at 6645.11 Å. They gave [Eu/Fe] = 0.51, which is much higher than the model predictions shown here, and also higher than Eu abundances for normal giant stars in Forsberg et al. (2019).

A.2.5. HD 59852 (Fig. A.8)

We show different models between 2.0 and 3.5 $M_{\odot}$ for this star. There are no error bars for three elements (Sr, Zr, Ru), and Nb is not available for this star. All models but the FRUITY T60 model match the *s*-process abundance pattern well, except Ru, which we thus kept in sub-group 1b. We conclude that an AGB with mass between 2 and 3 $M_{\odot}$ was the polluter for this Ba star, in agreement with M_{AGB}^{J19} .

A.2.6. HD 95193 (Fig. A.9)

We show FRUITY models between 2.0 and 3.0 $M_{\odot}$ and Monash models, with different ^{13}C pocket sizes, between 2.5 and 3.5 $M_{\odot}$. The FRUITY models with supersolar metallicity ([Fe/H] = 0.15), still within the error bar of the Ba star’s metallicity) overestimate the determined Sr, Y, and Zr abundances compared to other models shown here, and these models also have larger δ values than other model predictions for this star. All of the other models match the determined abundances well, although the [Rb/Fe] value is lower than most of the model predictions, indicating a lower than 3 $M_{\odot}$ initial AGB mass.

A.2.7. HD 119185 (Fig. A.10)

The abundance pattern of HD 119185 can be matched by low-mass (1.5–2.1 $M_{\odot}$) models very well. The [Rb/Fe] value is lower than solar, but within the error bar its value is in agreement with a low-mass (1.5–2.1 $M_{\odot}$) polluting AGB. As for HD 59852, Nb is absent and [Ru/Fe] is higher than the model predictions.

A.2.8. HD 143899 (Fig. A.11)

We show models between 2 and 3.25 $M_{\odot}$ with different ^{13}C pockets. The observed [La/Fe] and [Rb/Fe] values are somewhat

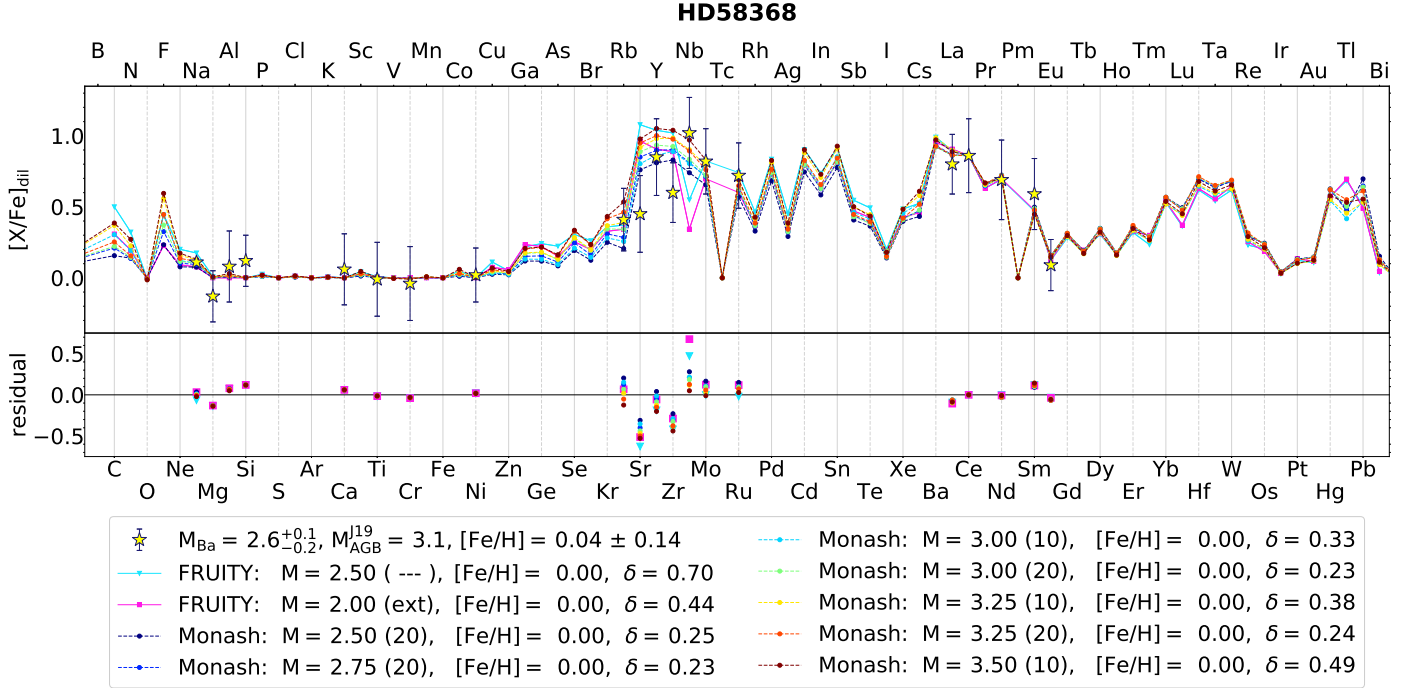


Fig. A.1. Same as Fig. 5, but for HD 58368.

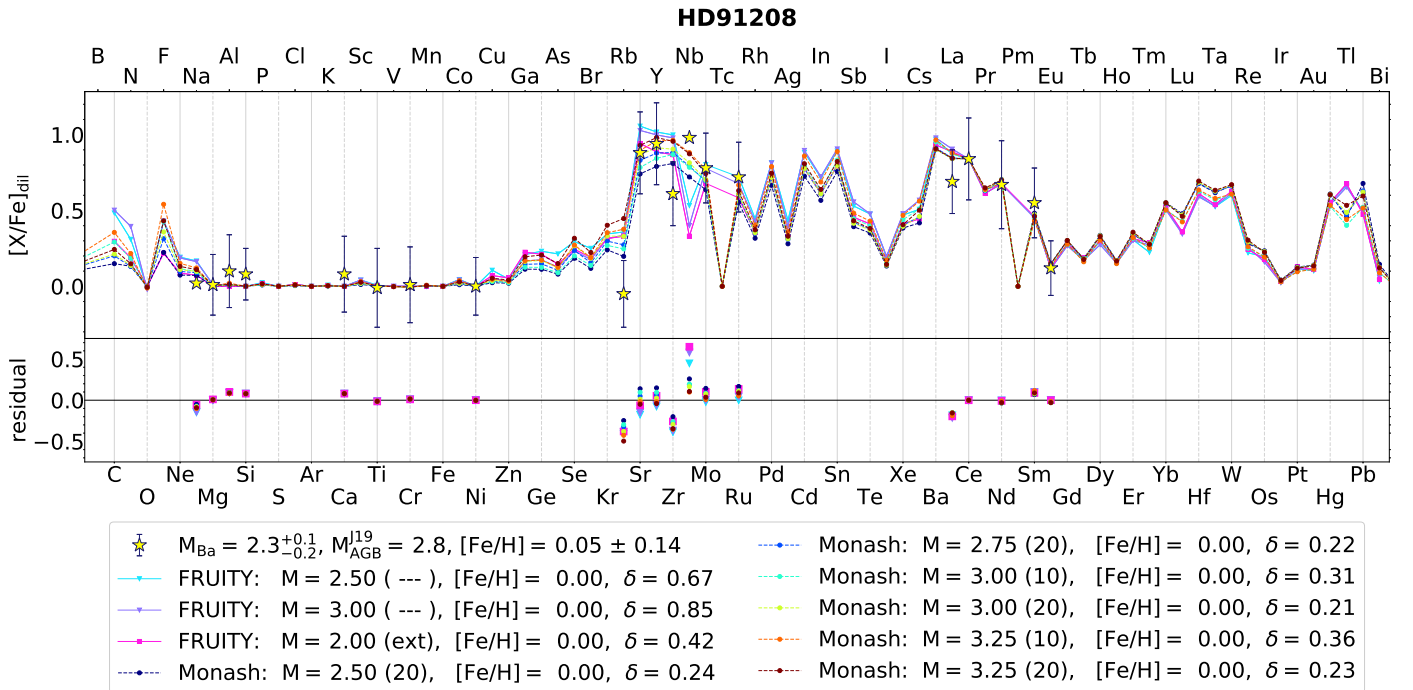


Fig. A.2. Same as Fig. 5, but for HD 91208.

lower and, as for the other stars in this sub-group, Nb is higher than the model predictions, but the error bar is lacking for this element due to the low number of lines used for the derivation of its abundance. We plot the $2.0 M_{\odot}$ FRUITY T60 model for completeness, because this model is at the limit of the metallicity range, but it highly overproduces the first peak elements compared to the observations, and the δ value of this model is much higher than for the others. Models between 2.5 and $3.0 M_{\odot}$ have the best match to the abundance pattern, which is in agreement with M_{AGB}^{19} of $3 M_{\odot}$.

A.2.9. HD 200063 (Fig. A.12)

We show different metallicity models between 1.9 and $2.75 M_{\odot}$. The uncertain mass of the Ba star and especially the strong assumption of the constant Q when deriving the WD mass could lead to a poorly estimated M_{AGB}^{19} . All of the models are within the error bar of $[Rb/Fe]$ and, as was in the case of HD 143899, the FRUITY T60 model largely overproduces the first peak elements. Higher metallicity FRUITY models and the highest mass Monash model with $2.75 M_{\odot}$ match the abundance pattern within

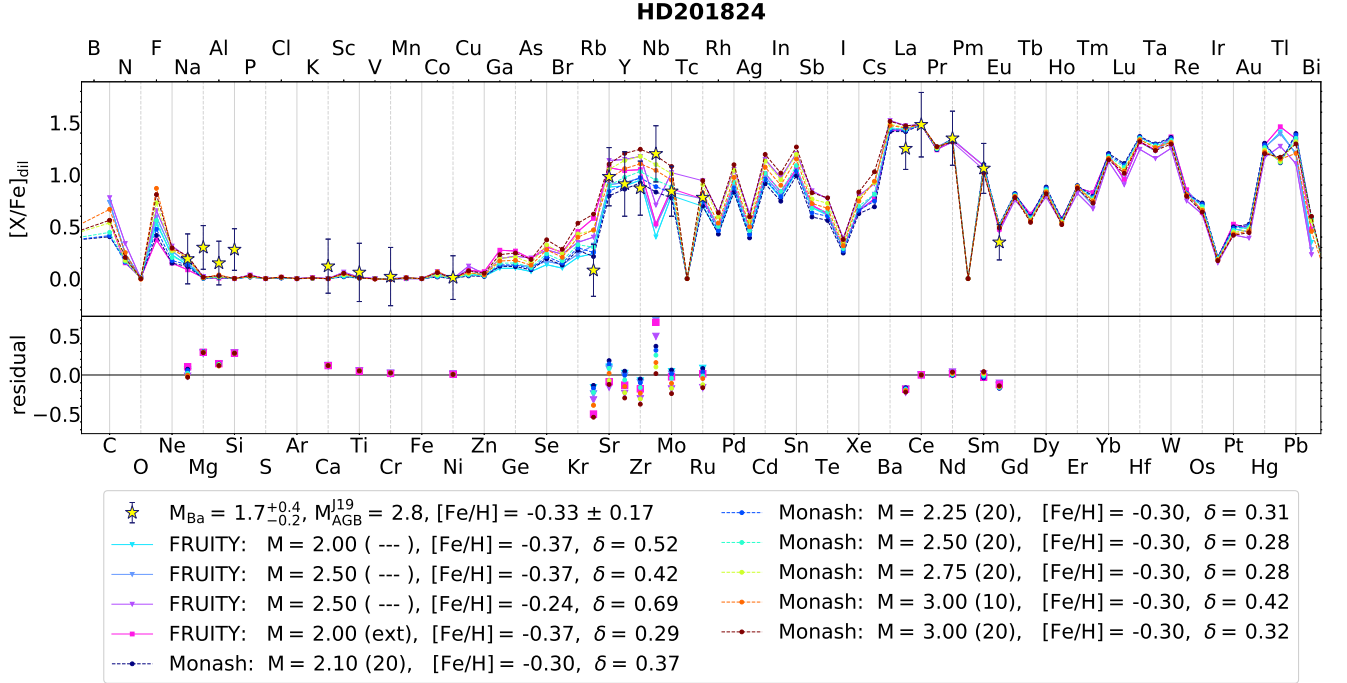


Fig. A.3. Same as Fig. 5, but for HD 201824.

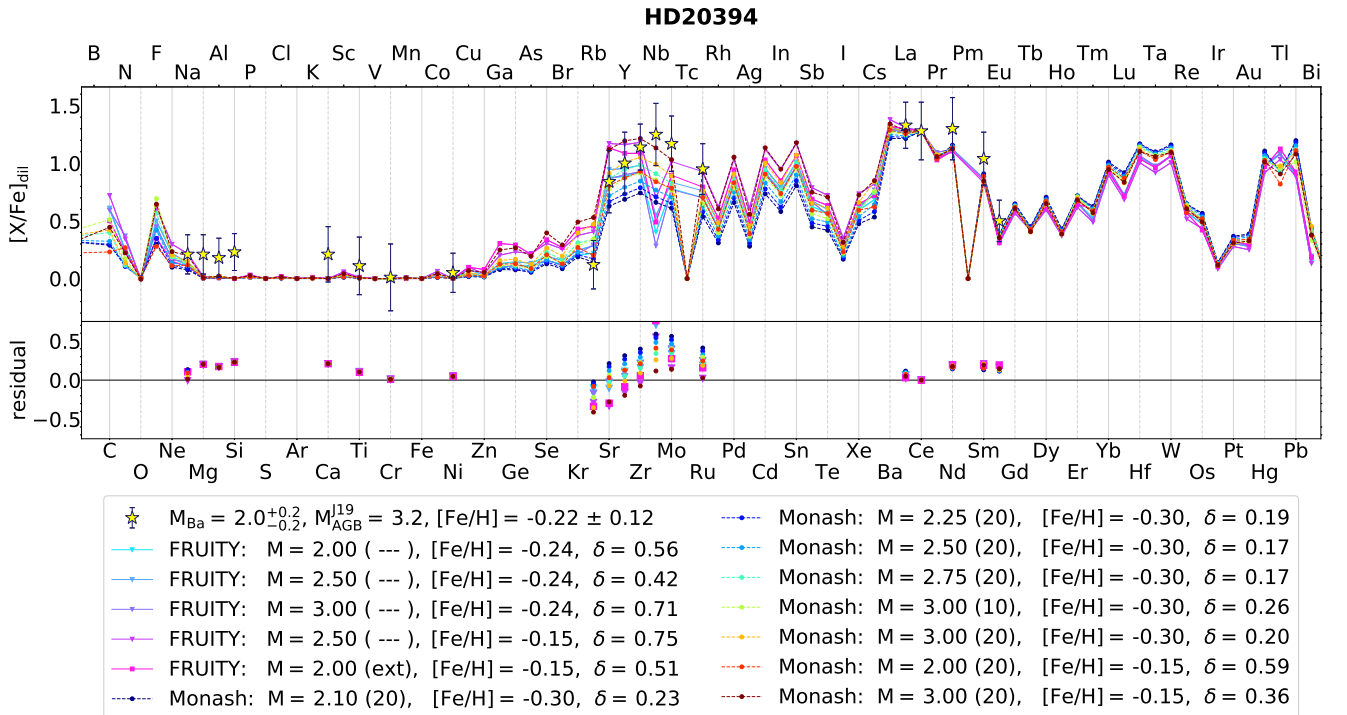


Fig. A.4. Same as Fig. 5, but for HD 20394.

the error bars. This mass range of the models is in agreement with M_{AGB}^{19} . Nb is higher than the model predictions, as for any other star in this sub-group.

A.2.10. HD 210946 (Fig. A.13)

A large variety of models with different metallicities between 1.5 and 2.25 $M_{\odot}$ are shown. The first *s*-process peak of this star shows higher abundances than the second peak elements. The

higher first peak elements along with Rb are in agreement with non-rotating solar metallicity models around 2 $M_{\odot}$. This is in agreement with M_{AGB}^{19} of 1.9 $M_{\odot}$. Nb is higher than the model predictions, as this star belongs to sub-group 1b.

A.2.11. HD 211594 (Fig. A.14)

We show FRUITY models with masses between 2.0 and 2.5 $M_{\odot}$ and Monash models between 2.5 and 3.5 $M_{\odot}$ for this star. Higher

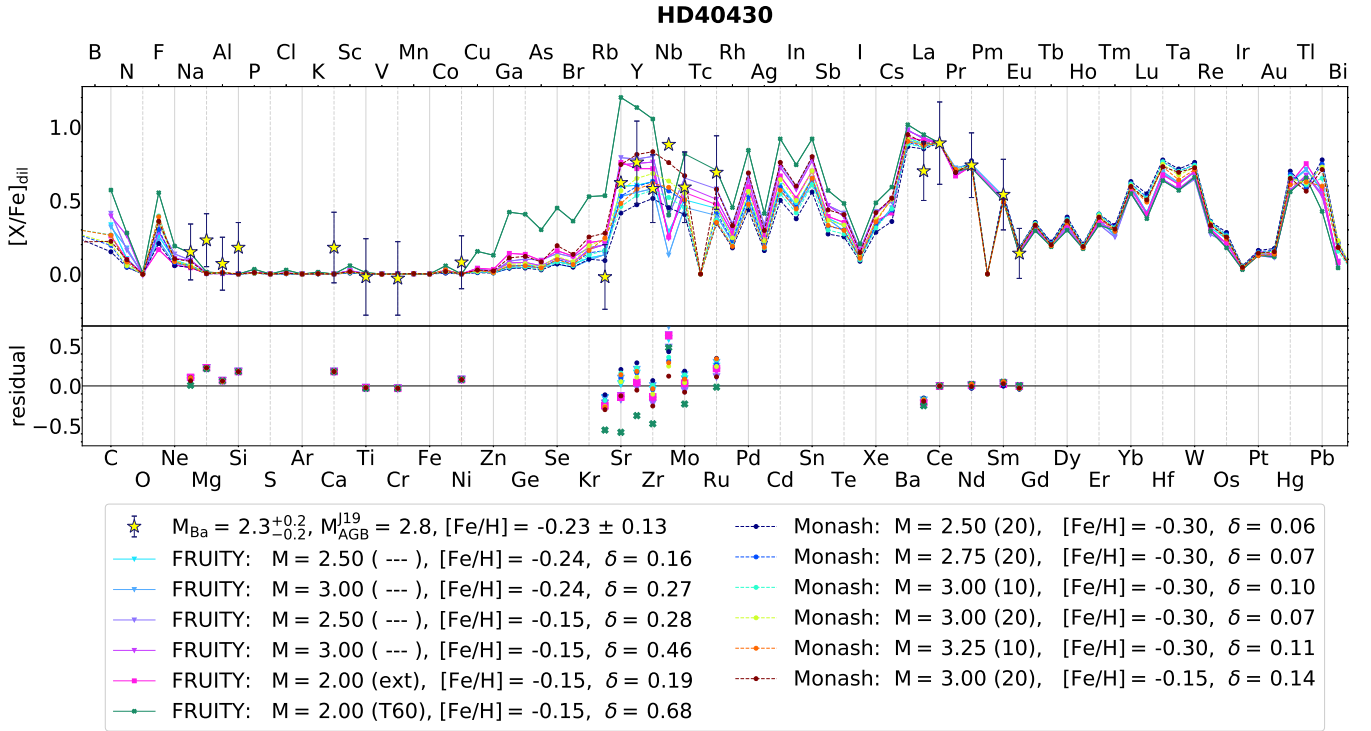


Fig. A.5. Same as Fig. 5, but for HD 40430.

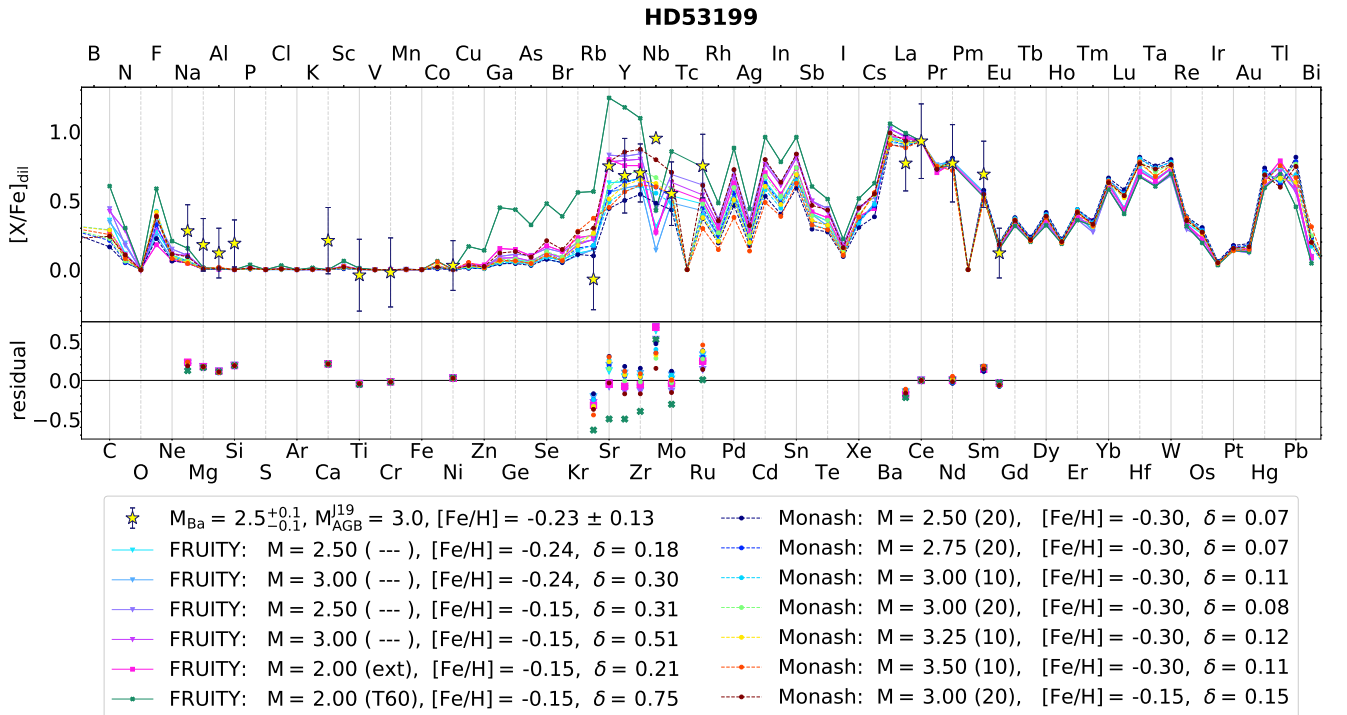


Fig. A.6. Same as Fig. 5, but for HD 53199.

mass FRUITY models have larger δ values than our limit of 0.9. The models match the *s*-process peaks well, except for Nb, which is higher than for the stars in sub-group 1b. The [Rb/Fe] value and the first peak elements point to a $\approx 3 M_{\odot}$ polluter AGB, in agreement with the estimated initial mass of $3.2 M_{\odot}$. We note that high δ values (larger than 0.3) are necessary to match the abundances' pattern.

A.3. Group 1c: Observed Nb, Mo, Ru, Nd, and Sm abundances higher than the model predictions

A.3.1. BD $-14^{\circ}2678$ (Fig. A.15)

The star BD $-14^{\circ}2678$ has M_{AGB}^{19} of $3.5 M_{\odot}$. The *s*-process peaks and light elements are in agreement with models in the mass range of $2-3.5 M_{\odot}$, although based on the determined Rb abun-

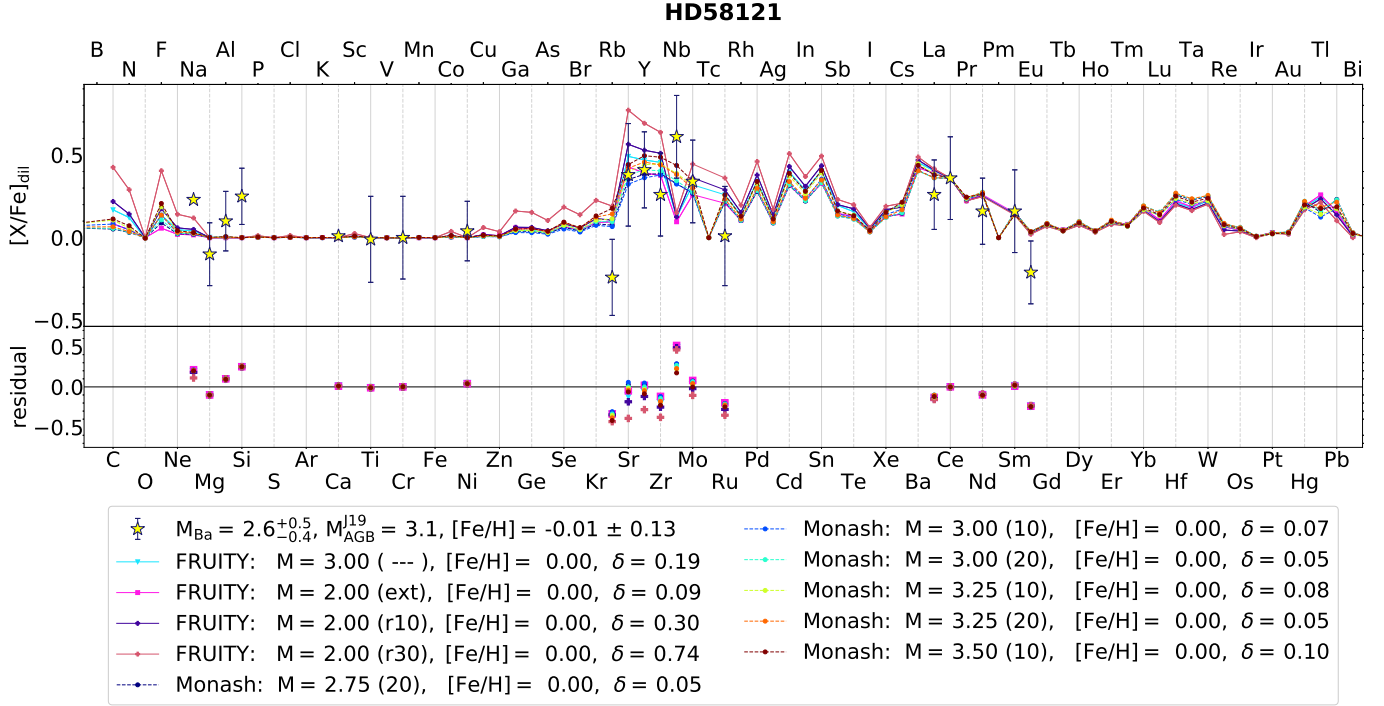


Fig. A.7. Same as Fig. 5, but for HD 58121.

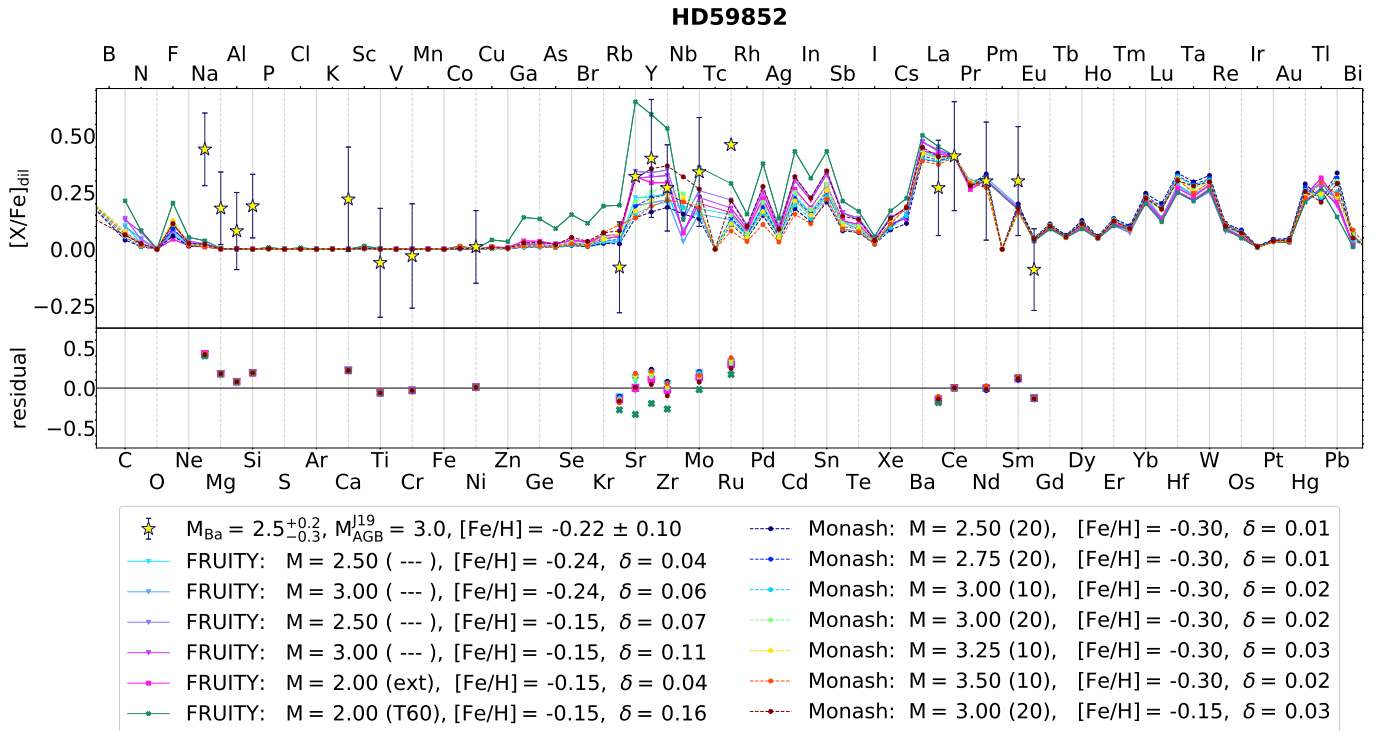


Fig. A.8. Same as Fig. 5, but for HD 59852.

dance the polluting AGB might be $\approx 3 M_{\odot}$ or lower rather than higher. Except from the $2.5 M_{\odot}$ model, all FRUITY models with standard pocket in the mass range of 2.0 – $4.0 M_{\odot}$ have higher δ values than our limit of $\delta = 0.9$. We also show $2 M_{\odot}$ FRUITY ‘TAIL’ and rotating ‘TAIL’ models for comparison, although these models overestimate the first peak abundances compared to the observations and thus are not a good match. As this star belongs to sub-group 1c, Nb, Mo, Ru, and Sm are somewhat

higher than the model predictions, although Nb and Ru have no error bar.

A.3.2. CPD $-64^{\circ}4333$ (Fig. A.16)

The star CPD $-64^{\circ}4333$ has a very short orbital period close to 1 yr. Since Gaia DR2 parallaxes are calculated from single-star solution, this could result in an unreliable mass of the Ba star, and

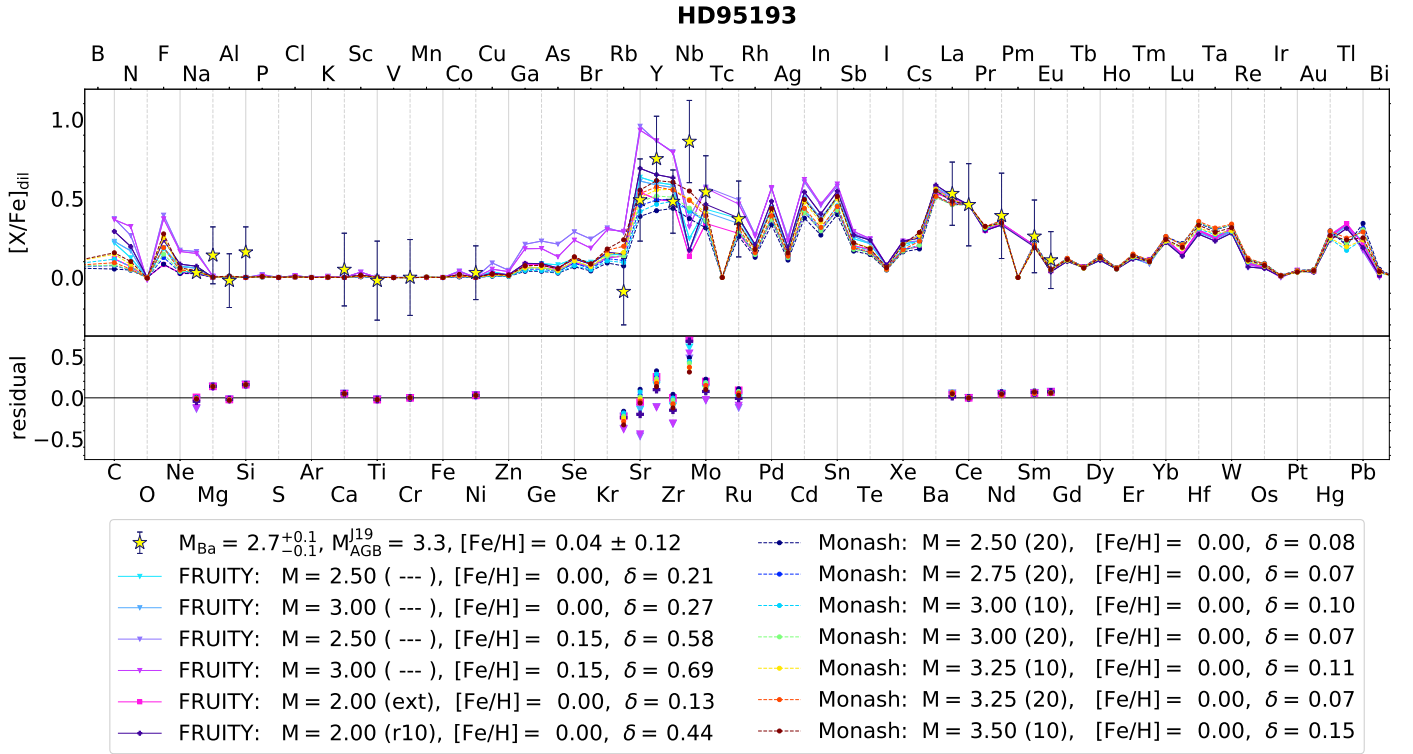


Fig. A.9. Same as Fig. 5, but for HD 95193.

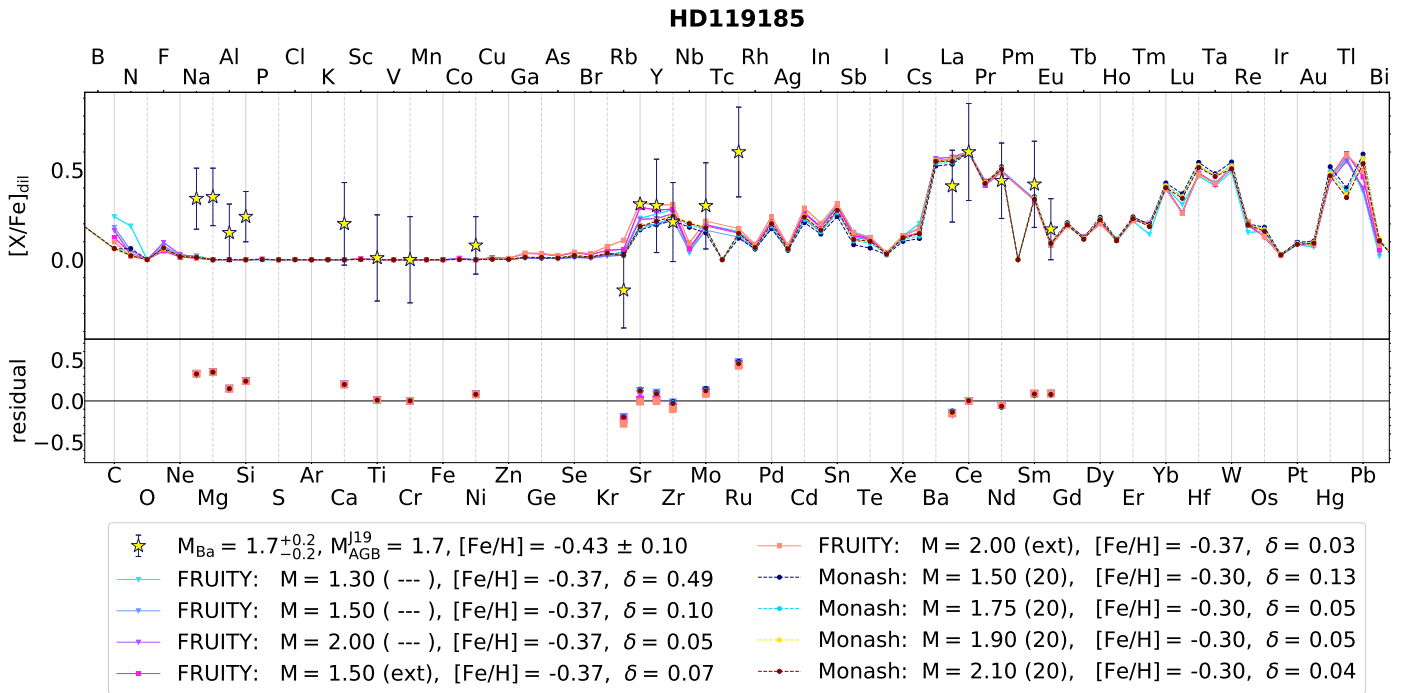


Fig. A.10. Same as Fig. 5, but for HD 119185.

thus M_{AGB}^{19} of $\approx 2 M_{\odot}$ can be very uncertain. This is supported by the value of the RUWE for this star of 4.181 (Gaia Collaboration 2020), indicating problems with the single-star astrometric solution due to the short orbital period. Here, we show different models with 2 and $2.5 M_{\odot}$. All of the models match the determined abundances well, in agreement with M_{AGB}^{19} .

However, the models show high δ values, which could be explained by the short period, and thus larger mass transfer efficiency in the system. The elements after the s -process peaks show higher values than the model predictions, as for other stars in sub-group 1c.

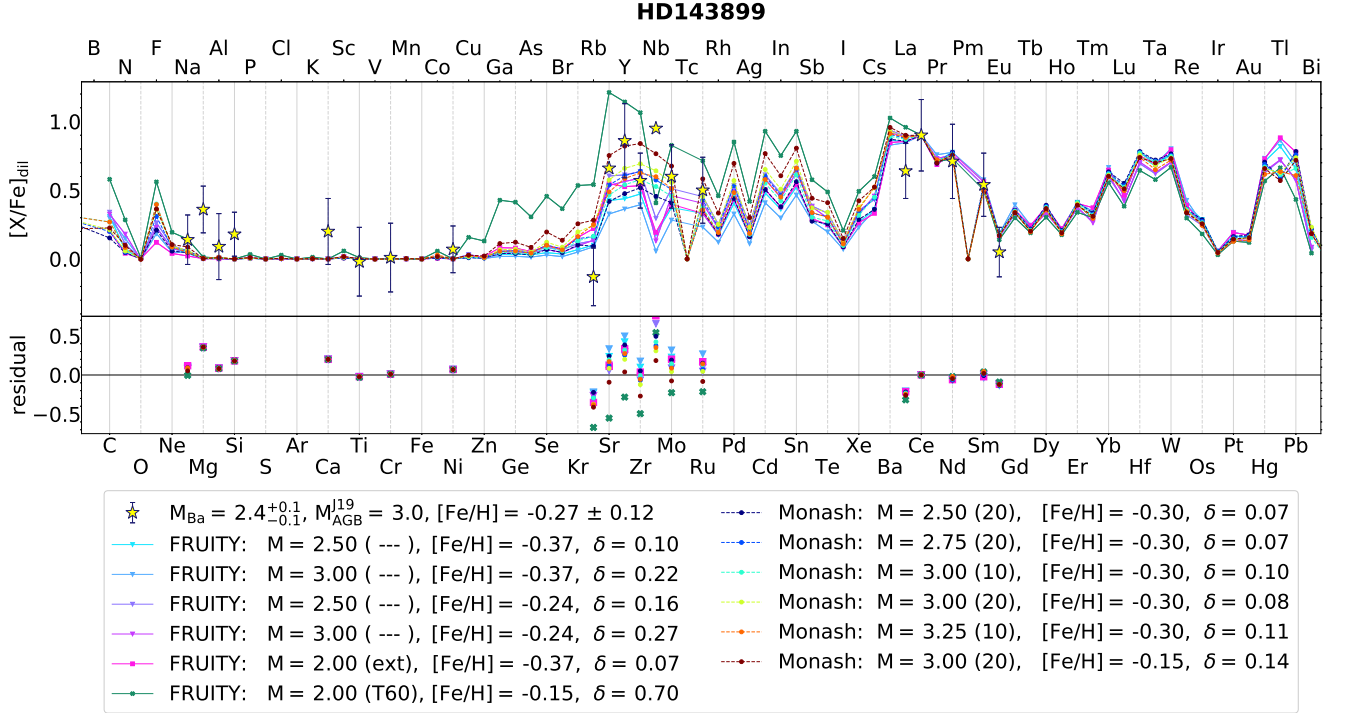


Fig. A.11. Same as Fig. 5, but for HD 143899.

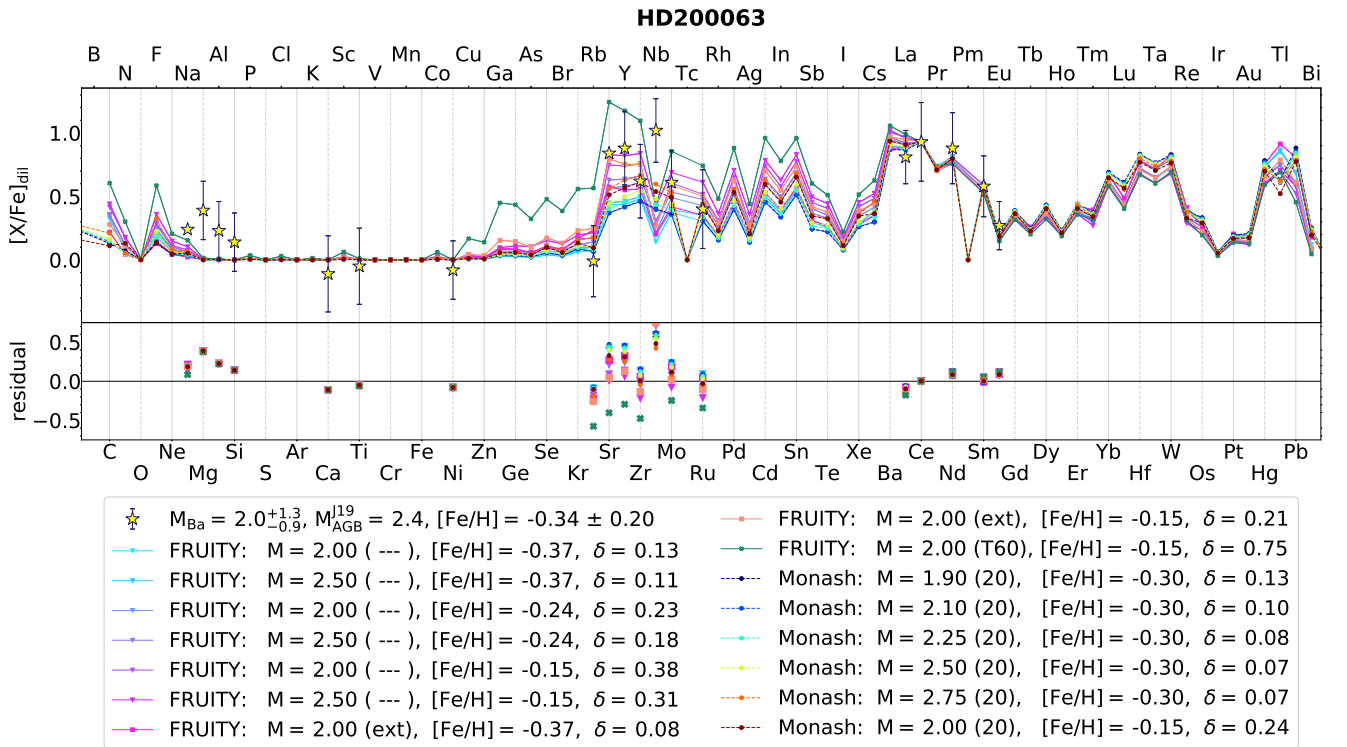


Fig. A.12. Same as Fig. 5, but for HD 200063.

A.3.3. HD 24035 (Fig. A.17)

The star HD 24035 has the shortest orbital period in our sample with 377.8 days. We show models between 1.5 and a $2.25 M_{\odot}$. Both of the *s*-process peaks are matched, but the elements after the peaks show higher abundances than the model predictions. The $[Rb/Fe]$ value suggests that an AGB star with $2 M_{\odot}$ was

the polluter in the system, in agreement with M_{AGB}^{19} . The models have high δ values, but for CPD $-64^{\circ}4333$, the short orbital period of the system could explain this phenomenon.

Shejeelammal et al. (2020) also analysed this star and found an initial AGB mass of $2.5 M_{\odot}$ based on χ^2 minimalisation using FRUITY models. They found $[Fe/H] = -0.5$ and generally higher *s*-process abundances than the data in this study.

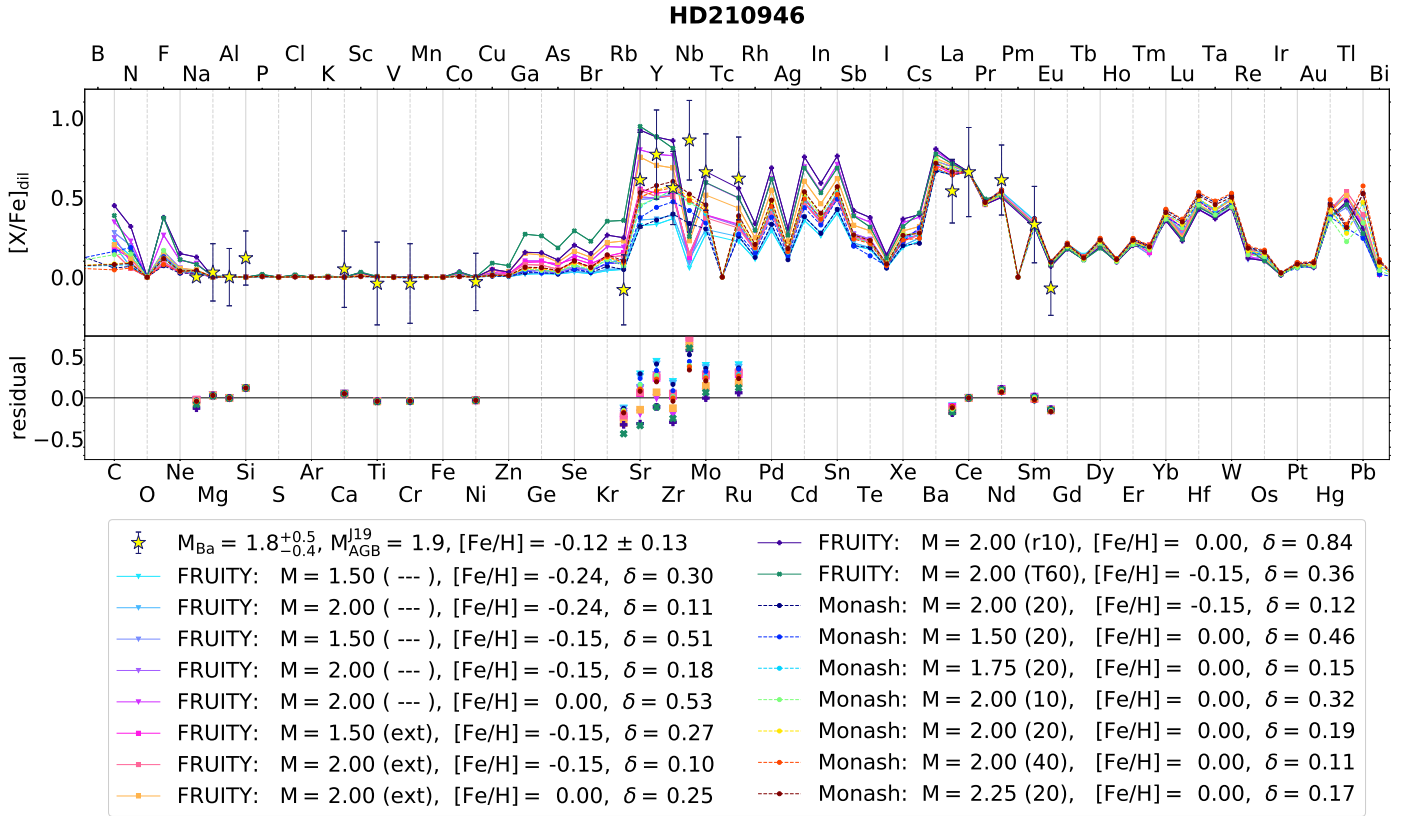


Fig. A.13. Same as Fig. 5, but for HD 210946.

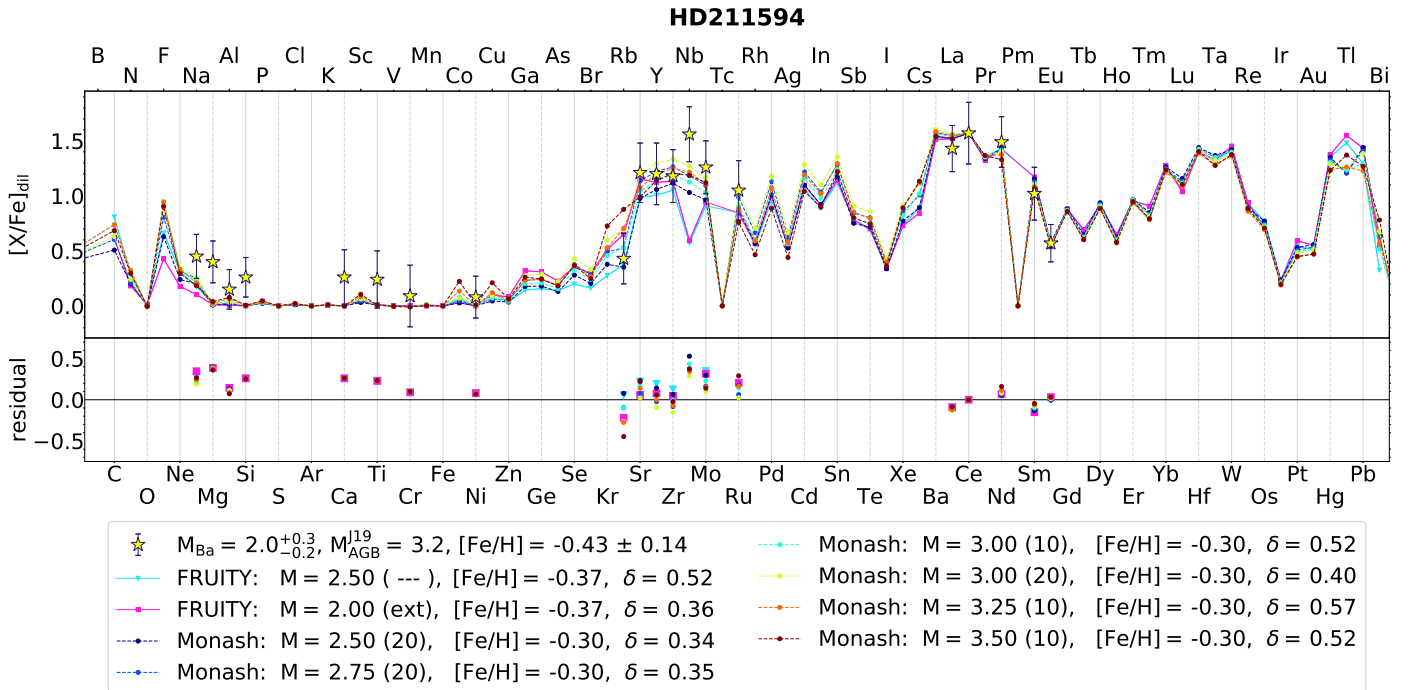


Fig. A.14. Same as Fig. 5, but for HD 211594.

A.3.4. HD 183915 (Fig. A.18)

Considering the large uncertainty of the Ba star mass, we plot models in the mass range between 1.5 and 2.5 $M_{\odot}$. All of these models are able to reproduce the abundance pattern at the

s-process peaks of this Ba star; however, the 2.0 $M_{\odot}$ FRUITY ‘TAIL’ model overestimates the [Rb/Fe] value compared to that of the star, and Nb, Mo, and Nd are higher than the model predictions.

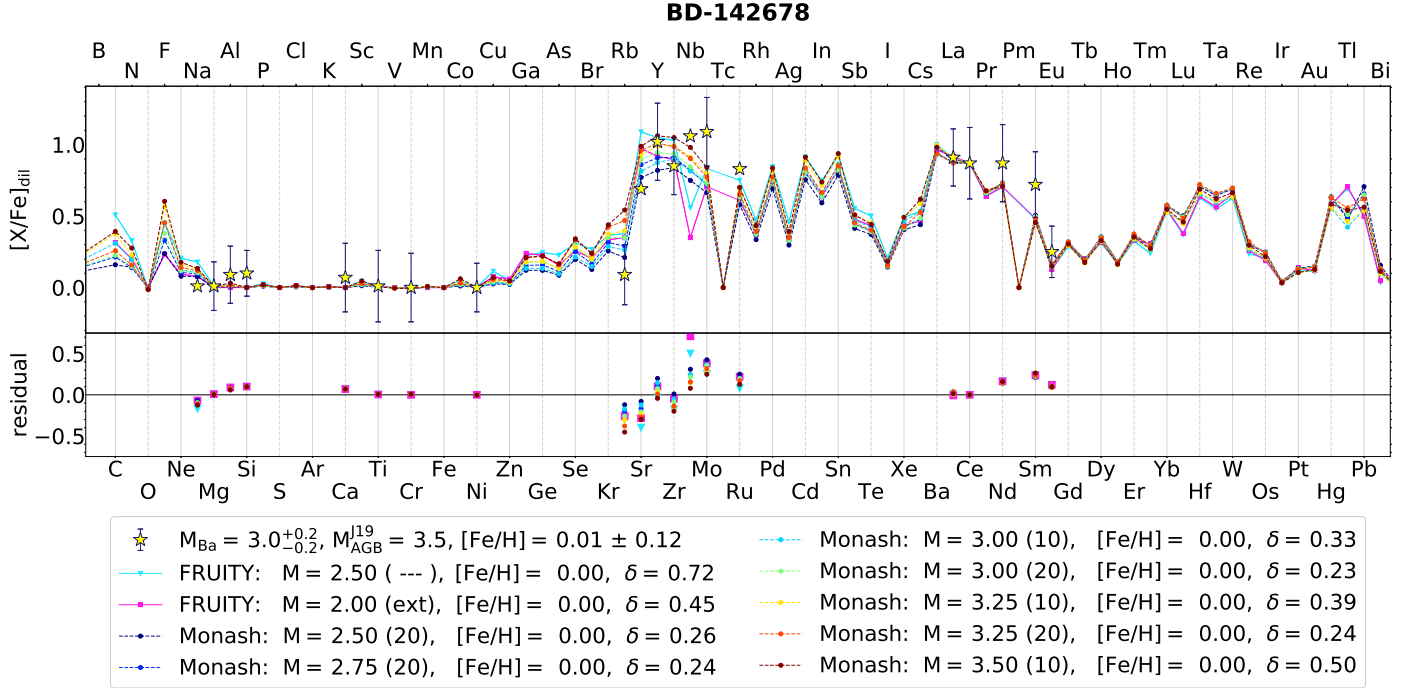


Fig. A.15. Same as Fig. 5, but for BD -14°2678.

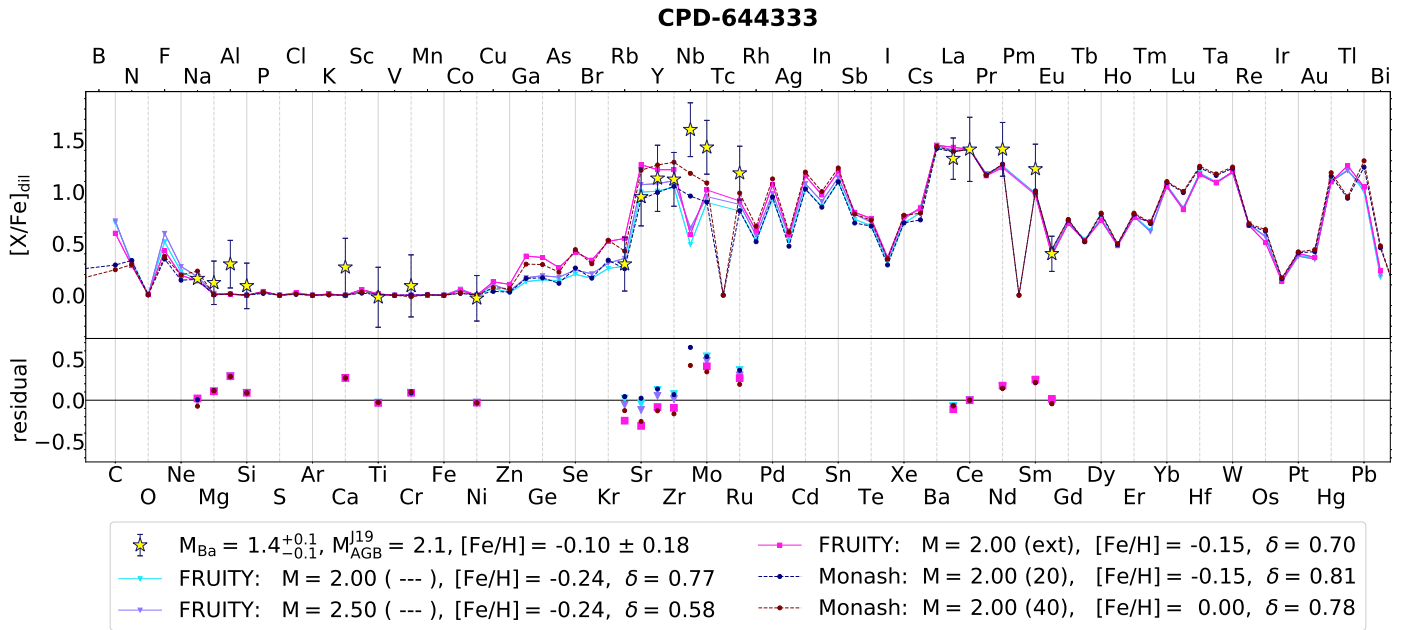


Fig. A.16. Same as Fig. 5, but for CPD -64°4333.

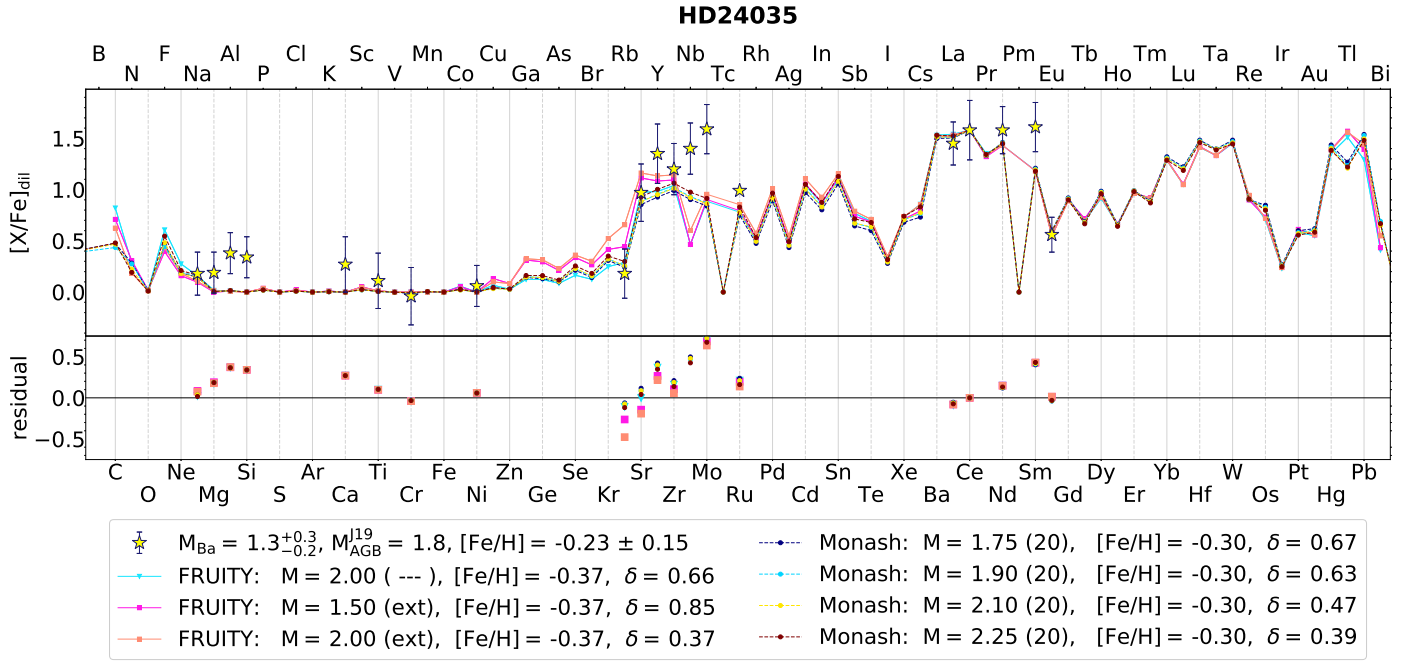


Fig. A.17. Same as Fig. 5, but for HD 24035.

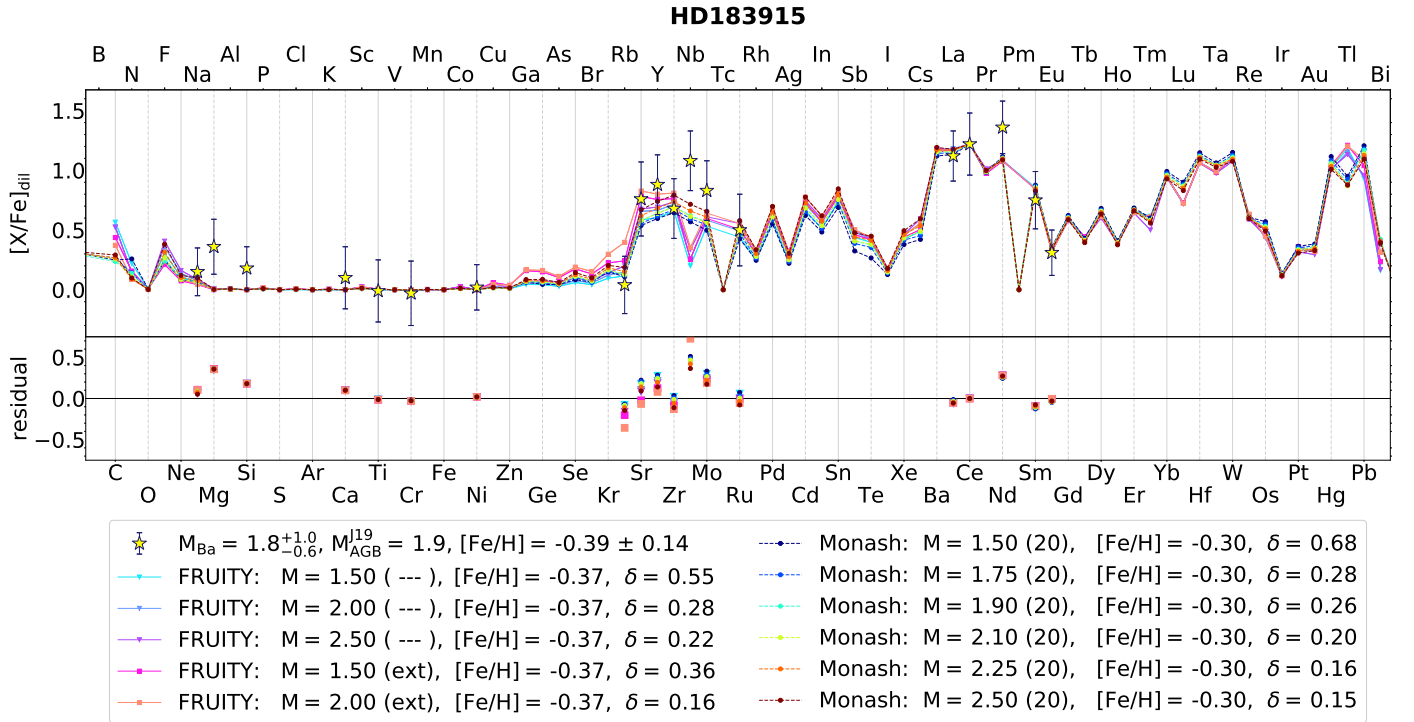


Fig. A.18. Same as Fig. 5, but for HD 183915.

Appendix B: Tables

All models with a ^{13}C pocket in the metallicity range between $[\text{Fe}/\text{H}] = -0.67$ and 0.15 and between -0.7 and 0 , and with mass

from 1.3 to 3 and from 1.25 to $4.5 M_{\odot}$ (FRUITY, Table B.1 and Monash, Table B.2, respectively) used for comparison for at least one of the Ba stars in the sample (see Sect. 3.1 for the description of the models).

Table B.1. FRUITY models used in this study.

M ($M_{\odot}$) \ $[\text{Fe}/\text{H}]$ (dex)	-0.67	-0.37	-0.24	-0.15	0.00	0.15
1.3	ST	ST				
1.5	ST, ext	ST, ext	ST	ST, ext, T60	ST	
2	ST, ext	ST, ext	ST	ST, ext, T60	ST, ext, r10, r30	
2.5		ST	ST	ST	ST	ST
3		ST	ST	ST	ST	ST

Notes. ST means the standard ^{13}C -pocket size, while ext, T60, r10, and r30 stand for ‘TAIL’, rotating ‘TAIL’ (60 km/s), and rotating models with 10 and 30 km/s initial rotation velocity, respectively.

Table B.2. Monash models used in this study.

M ($M_{\odot}$) \ $[\text{Fe}/\text{H}]$ (dex)	-0.70	-0.30	-0.15	0.00
1.25	20, 60			
1.50	20, 60	20		20
1.75	20	20		20
1.90		20		
2.00	20, 60		20	10, 20, 40
2.10		20		
2.25		20		20
2.50		20		20
2.75		20		20
3.00		10, 20	20	10, 20
3.25		10		10, 20
3.50		10		10
3.75		10		
4.00		1, 10		
4.25				10
4.50				10

Notes. The ^{13}C pocket size is given in the cells in units of $10^{-4} M_{\odot}$ for each model.